



PUBLISHED FOR SISSA BY SPRINGER

RECEIVED: February 3, 2016

ACCEPTED: April 1, 2016

PUBLISHED: April 11, 2016

Rainbow valley of colored (anti) de Sitter gravity in three dimensions

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ABSTRACT: We propose a theory of three-dimensional (anti) de Sitter gravity carrying Chan-Paton color charges. We define the theory by Chern-Simons formulation with the gauge algebra $(\mathfrak{gl}_2 \oplus \mathfrak{gl}_2) \otimes \mathfrak{u}(N)$, obtaining a color-decorated version of interacting spin-one and spin-two fields. We also describe the theory in metric formulation and show that, among N^2 massless spin-two fields, only the singlet one plays the role of metric graviton whereas the rest behave as *colored spinning matter* that strongly interacts at large N . Remarkably, these *colored spinning matter* acts as Higgs field and generates a non-trivial potential of staircase shape. At each extremum labelled by $k = 0, \dots, [\frac{N-1}{2}]$, the $\mathfrak{u}(N)$ color gauge symmetry is spontaneously broken down to $\mathfrak{u}(N-k) \oplus \mathfrak{u}(k)$ and provides different (A)dS backgrounds with the cosmological constants $(\frac{N}{N-2k})^2 \Lambda$. When this symmetry breaking takes place, the spin-two Goldstone modes combine with (or are eaten by) the spin-one gauge fields to become partially-massless spin-two fields. We discuss various aspects of this theory and highlight physical implications.

KEYWORDS: Space-Time Symmetries, Classical Theories of Gravity, Gauge-gravity correspondence

ARXIV EPRINT: [1511.05220](https://arxiv.org/abs/1511.05220)

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*“It’s time to try. Defying gravity.
I think I’ll try. Defying gravity.
And you can’t pull me down!”*

— Gregory Maguire

‘Wicked: The Life and Times of the Wicked Witch of the West’

1 Introduction

The Einstein’s theory of gravity is known to be rigid. Variety of modifications has been challenged with diverse motivations, yet no concrete result of success has been reported so far (for related readings, see e.g. [1–6] and references therein). Recently, two situations defying the rigidity of Einstein gravity were actively explored. One is the massive modification of gravity [7–10], along with numerous variants in three dimensions [11, 12]. Another is higher-derivative modifications of the gravity [13, 14].

In this work, we investigate the modification of Einstein gravity to a multi-graviton theory: *the color decoration*. In spite of previous negative results [15–18], certain models of colored gravity can be consistently constructed by introducing other field contents than massless spin-two fields only. Moreover, the color-decoration we study is not limited to the Einstein gravity and can be applied to various extensions of it. In particular, all higher-spin theories formulated in [19–21] can be straightforwardly color-decorated, whose first steps were conceived in [22–26]. In the companion paper [27], we study a three-dimensional color-decorated higher-spin gravity.

The color decoration of gravity evokes various conceptual issues. Clearly, the colored gravity is analogous to Yang-Mills theory were if the Einstein gravity compared to Maxwell theory. Besides the presence of multiple gauge bosons in the system, the Yang-Mills theory as color-decorated Maxwell theory has far-reaching consequences that are not shared by the Maxwell theory.¹ Likewise, we anticipate that color-decorated gravity brings out surprising new features one could not simply guess on a first look. In this paper, we define and study a version of the color-decorated Einstein gravity in three dimensions, and uncover remarkable new features not shared by the Einstein gravity itself. Most interestingly, we will find that this color-decorated gravity admits a number of (A)dS backgrounds with different cosmological constants as classical vacua.

In analyzing our model of three-dimensional color-decorated gravity, we shall make use of both the Chern-Simons formulation [28, 29] and the metric formulation. Various features of the theory are more transparent in one formulation over the other. For instance, the existence of multiple (A)dS vacua with different cosmological constants can be understood more intuitively in the metric formulation, whereas consistency of the theory is more manifest in the Chern-Simons formulation. The latter makes use of the gauge algebra,

$$\mathfrak{g} = (\mathfrak{gl}_2 \oplus \mathfrak{gl}_2) \otimes \mathfrak{u}(N), \quad (1.1)$$

where the $\mathfrak{u}(N)$ and $\mathfrak{gl}_2 \oplus \mathfrak{gl}_2$ correspond respectively to the color gauge algebra and the extended isometry algebra governing the gravitational dynamics. We stress that, compared to the usual gravity with $\mathfrak{sl}_2 \oplus \mathfrak{sl}_2$ gauge algebra, the color-decorated gravity has two additional identity generators from each of \mathfrak{gl}_2 . They are indispensable for the consistency of color decoration and correspond to two additional Chern-Simons gauge fields on top of the graviton. Hence, when colored-decorated, we get a massless spin-two field and two non-Abelian spin-one fields, both taking adjoint values of $\mathfrak{u}(N)$. Let us also remark that compared to the spin-one situation where the Abelian Maxwell theory turns into the non-Abelian Yang-Mills theory once color-decorated, Einstein gravity is already non-Abelian, while color decoration enlarges the gauge algebra of the theory.

Re-expressing the theory in metric formulation makes it clear that, among N^2 massless spin-two fields, only the singlet one plays the role of genuine graviton, viz. the first fundamental form, whereas the rest rather behave as *colored spinning matter* fields with

¹The story goes that, during C.N. Yang’s seminar at the Institute for Advanced Study at Princeton in 1953, Wolfgang Pauli commented that he first discovered non-Abelian gauge theory in this manner, but then immediately dismissed it because vector bosons are massless and hence “unphysical”. We acknowledge Stanley Deser for straightening us up for details of this history.

minimal covariant coupling to the gravity as well as to the $u(N)$ gauge fields. We derive the explicit form of Lagrangian for these *colored spinning matter* fields and find that they have a strong self-coupling compared to the gravitational one by the factor of \sqrt{N} . Analyzing the potential of the Lagrangian, we identify all the extrema: there are $[\frac{N+1}{2}]$ number of them and they have different cosmological constants,

$$\left(\frac{N}{N-2k}\right)^2 \Lambda, \quad (1.2)$$

where $k = 0, \dots, [\frac{N-1}{2}]$ is the label of the extrema and Λ is the cosmological constant of the vacuum with maximum radius (corresponding to $k = 0$). Note that not only (A)dS but also any exact gravitational backgrounds such as BTZ black holes [30] lie multiple times with different cosmological constants (1.2) in the vacua of the colored gravity. All extrema except the $k = 0$ vacuum spontaneously break the color symmetry $U(N)$ down to $U(N-k) \times U(k)$. When this symmetry breaking takes place, the corresponding $2k(N-k)$ spin-two Goldstone modes are combined with the gauge fields to become the partially-massless spin-two fields [31, 32]: the latter spectrum does not have any propagating degrees of freedom (DoF) similarly to the massless ones. Instead in AdS case, they describe ‘four’ boundary DoF which originate from the boundary modes of the colored massless spin-two and spin-one fields.

The organization of the paper is as follows. In section 2, we recapitulate the no-go theorem of interacting theory of multiple massless spin-two fields. In section 3, we define the color-decorated (A)dS₃ gravity in Chern-Simons formulation, and discuss how this theory evades the no-go theorem. In section 4, we recast the Chern-Simons action into metric formulation by solving torsion condition and obtain the Lagrangian for the colored massless spin-two fields. In section 5, we solve the equations of motion and find a class of classical vacua with varying degrees of color symmetry breaking. We show that these (A)dS vacua have different cosmological constants. We explicitly investigate the simplest example of $k = 1$ vacuum in $N = 3$ case. In section 6, we expand the theory around a color non-singlet vacuum and analyze the field spectrum contents. We demonstrate that the fields corresponding to the broken part of the color symmetry describe the spectrum of partially-massless spin-two field. Section 7 contains discussions of our results and outlooks. Finally, appendix A reviews massive and (partially-)massless spin-two fields in three-dimensions.

2 No-go theorem on multiple spin-two theory

Einstein gravity describes the dynamics of massless spin-two field on a chosen vacuum. Conversely, it can also be verified that the Einstein gravity is the only interacting theory of a massless spin-two field (see e.g. [33–37]). In this context, one may ask whether there exists a non-trivial theory of multiple massless spin two fields. This possibility has been examined in [15–18], leading to a no-go theorem. We shall begin our discussion by reviewing this result.²

²See also related discussion in [38, 39].

The no-go theorem asserts that there is no interacting theory of multiple massless spin-two fields, without inclusion of other fields. The first point to note in this consideration is that any gauge-invariant two-derivative cubic interactions among the spin-two fields is in fact equivalent to that of Einstein-Hilbert (EH) action, modulo *color-decorated* cubic coupling constants g_{IJK} :

$$g_{IJK} \left(h_{\mu\rho}^I \partial^\rho h_{\nu\lambda}^J \partial^\lambda h^{K\mu\nu} + \dots \right). \quad (2.1)$$

Here, $h_{\mu\nu}^I$ are the massless spin-two fields with *color* index I , and the tensor structure inside of the bracket is that of the EH cubic vertex. For the consistency with the color indices, it is required that the coupling constants are fully symmetric: $g_{IJK} = g_{(IJK)}$. Moreover, the gauge invariance requires that these constants define a Lie algebra spanned by the colored isometry generators. For instance, in the Minkowski spacetime, the colored generators P_μ^I and $M_{\mu\nu}^I$ obey

$$[M_{\mu\nu}^I, P_\rho^J] = 2 g^{IJ}{}_K \eta_{\rho[\nu} P_{\mu]}^K, \quad [M_{\mu\nu}^I, M_{\rho\lambda}^J] = 4 g^{IJ}{}_K \eta_{[\nu[\rho} M_{\lambda]\mu]}^K. \quad (2.2)$$

Relating these colored generators to the usual isometry ones as $P_\mu^I = P_\mu \otimes \mathbf{T}^I$ and $M_{\mu\nu}^I = M_{\mu\nu} \otimes \mathbf{T}^I$, one can straightforwardly conclude that the color algebra \mathfrak{g}_c generated by \mathbf{T}^I must be *commutative* and *associative* [15–17]. Moreover, one can even show that \mathfrak{g}_c necessarily reduces to a direct sum of one-dimensional ideals [18]: $\mathbf{T}^I \mathbf{T}^J = 0$ for $I \neq J$. Therefore, in this set-up, the only possibility is the simple sum of several copies of Einstein gravity which do not interact with each other.

This no-go theorem can be evaded with a slight generalization of the setup. Firstly, if the isometry algebra can be consistently extended from a Lie algebra to an associative one, then the commutativity condition on the color algebra \mathfrak{g}_c can be relaxed. The associative extension of isometry algebra typically requires to include other spectra, such as spin-one and possibly higher spins [22–26]. Moreover, it is not necessary to require that the structure constants g_{IJK} of \mathfrak{g}_c be totally symmetric, but sufficient to assume that the totally symmetric part is non-vanishing, $g_{(IJK)} \neq 0$, so that massless spin-two fields have non-trivial interactions among themselves.

Hence, an interacting theory of multiple massless spin-two fields might be viable once other fields are added and coupled to them. As the next consistency check, one can examine the fate of the general covariance in such a theory: if there exists a genuine metric field among these massless spin-two fields, the others should be subject to interact covariantly with gravity. Moreover, one can also examine whether the multiple massless spin two fields can be color-decorated bona fide by carrying non-Abelian charges. In principle, a theory can be made to covariantly interact with gravity or non-Abelian gauge field by simply replacing all its derivatives by the covariant ones with respect to both the diffeomorphism transformation and the non-Abelian gauge transformation. However, as in the diffeomorphism-covariant interactions of higher-spin fields, such replacements spoil the gauge invariance of the original system [40, 41]. The problematic term in the gauge variation is proportional to the curvatures, namely, Riemann tensor $R_{\mu\nu\rho\lambda}$ or non-Abelian gauge field strength $F_{\mu\nu}$. In three-dimensions, fortuitously, this is not a problem as these

curvatures are just proportional to the field equations of Einstein gravity or Chern-Simons gauge theory, respectively. In higher dimensions, these terms can be compensated by introducing a non-trivial cosmological constant, but at the price of adding higher-derivative interactions [42–44].

So, we conclude that, to have a consistent interacting theory of color-decorated massless spin-two fields, we need an (A)dS isometry gauge algebra which can be extended to an associative one. An immediate candidate is higher-spin algebra in any dimensions, since Vasiliev’s higher-spin theory can be consistently color-decorated, as mentioned before. Other option is to take the isometry algebras of (A)dS₃ and (A)dS₅ which are isomorphic to $\mathfrak{sl}_2 \oplus \mathfrak{sl}_2$ and \mathfrak{sl}_4 and can be extended to associative ones, $\mathfrak{gl}_2 \oplus \mathfrak{gl}_2$ and \mathfrak{gl}_4 by simply adding unit elements corresponding to spin-one fields.

3 Color-decorated (A)dS₃ gravity: Chern-Simons formulation

Let us now move to the explicit construction of a theory of colored gravity. In this paper, we focus on the case of three-dimensional gravity.

3.1 Color-decorated Chern-Simons gravity

In the uncolored case, it is known that the three-dimensional gravity can be formulated as a Chern-Simons theory with the action

$$S[\mathcal{A}] = \frac{\kappa}{4\pi} \int \text{Tr} \left(\mathcal{A} \wedge d\mathcal{A} + \frac{2}{3} \mathcal{A} \wedge \mathcal{A} \wedge \mathcal{A} \right), \quad (3.1)$$

for the gauge algebra $\mathfrak{sl}_2 \oplus \mathfrak{sl}_2$. The constant κ is the level of Chern-Simons action. We are interested in color-decorating this theory. Physically, this can be done by attaching Chan-Paton factors to the gravitons. Mathematically, this amounts to requiring the fields to take values in the tensor-product space $\mathfrak{g}_i \otimes \mathfrak{g}_c$, where the \mathfrak{g}_i is the isometry part of the algebra including $\mathfrak{sl}_2 \oplus \mathfrak{sl}_2$ and the \mathfrak{g}_c is a finite-dimensional Lie algebra of a matrix group \mathfrak{G}_c . For generic Lie algebras \mathfrak{g}_i and \mathfrak{g}_c , their tensor product do not form a Lie algebra, as is clear from the commutation relations:

$$[M_X \otimes \mathbf{T}_I, M_Y \otimes \mathbf{T}_J] = \frac{1}{2} [M_X, M_Y] \otimes \{\mathbf{T}_I, \mathbf{T}_J\} + \frac{1}{2} \{M_X, M_Y\} \otimes [\mathbf{T}_I, \mathbf{T}_J]. \quad (3.2)$$

The anticommutators $\{\mathbf{T}_I, \mathbf{T}_J\}$ and $\{M_X, M_Y\}$ do not make sense within the Lie algebras. Instead, if we start from associative algebras \mathfrak{g}_i and \mathfrak{g}_c , their direct product $\mathfrak{g}_i \otimes \mathfrak{g}_c$ will form an associative algebra, from which we can also obtain the Lie algebra structure. Hence, in this paper, we will consider associative algebras for \mathfrak{g}_i and \mathfrak{g}_c . For the color algebra \mathfrak{g}_c , we take the matrix algebra $\mathfrak{u}(N)$. For the isometry algebra \mathfrak{g}_i , we take $\mathfrak{g}_i = \mathfrak{gl}_2 \oplus \mathfrak{gl}_2$ (instead of $\mathfrak{sl}_2 \oplus \mathfrak{sl}_2$). The trace Tr of (3.1) should be defined also in the tensor product space and is given by the product of two traces as

$$\text{Tr}(\mathfrak{g}_i \otimes \mathfrak{g}_c) := \text{Tr}(\mathfrak{g}_i) \text{Tr}(\mathfrak{g}_c). \quad (3.3)$$

We also need for the fields to obey Hermiticity conditions compatible with the real form of the complex algebra.³

Therefore, our model of colored gravity is the Chern-Simons theory (3.1) where the one-form gauge field \mathcal{A} takes value in

$$\mathfrak{g} = (\mathfrak{gl}_2 \oplus \mathfrak{gl}_2) \otimes \mathfrak{u}(N) \ominus \text{id} \otimes \mathbf{I}. \quad (3.4)$$

Notice that we have subtracted the $\text{id} \otimes \mathbf{I}$, where id and \mathbf{I} are the centers of $\mathfrak{gl}_2 \oplus \mathfrak{gl}_2$ and $\mathfrak{u}(N)$, respectively: it corresponds to an Abelian vector field (described by Chern-Simons action) which does not interact with other fields.⁴ As a complex Lie algebra, \mathfrak{g} in (3.4) is in fact isomorphic to $\mathfrak{sl}_{2N} \oplus \mathfrak{sl}_{2N}$. This can be understood from the fact that the tensor product of 2×2 and $N \times N$ matrices gives $2N \times 2N$ matrix. It would be worth to remark as well that the algebra \mathfrak{g} necessarily contains elements in $\text{id} \otimes \mathfrak{su}(N)$ which correspond to the gauge symmetries of $\mathfrak{su}(N)$ Chern-Simons theory. In this sense, this $\mathfrak{su}(N)$ will be referred to as the color algebra.

It turns out useful⁵ to decompose the algebra \mathfrak{g} (3.4) into two orthogonal parts as

$$\mathfrak{g} = \mathfrak{b} \oplus \mathfrak{c}, \quad \text{such that} \quad \text{Tr}(\mathfrak{b} \mathfrak{c}) = 0, \quad (3.5)$$

where \mathfrak{b} is the subalgebra:

$$[\mathfrak{b}, \mathfrak{b}] \subset \mathfrak{b}, \quad (3.6)$$

corresponding to the *gravity plus gauge* sector (mediating gravity and gauge forces), whereas \mathfrak{c} corresponds to the *matter* sector — including all colored spin-two fields — subject to the covariant transformation,

$$[\mathfrak{b}, \mathfrak{c}] \subset \mathfrak{c}. \quad (3.7)$$

Corresponding to the decomposition (3.5), the one-form gauge field \mathcal{A} can be written as the sum of two parts

$$\mathcal{A} = \mathcal{B} + \mathcal{C}, \quad (3.8)$$

where \mathcal{B} and \mathcal{C} takes value in \mathfrak{b} and \mathfrak{c} , respectively. In terms of \mathcal{B} and \mathcal{C} , the Chern-Simons action (3.1) is reduced to

$$S[\mathcal{B}, \mathcal{C}] = \frac{\kappa}{4\pi} \int \text{Tr} \left(\mathcal{B} \wedge d\mathcal{B} + \frac{2}{3} \mathcal{B} \wedge \mathcal{B} \wedge \mathcal{B} + \mathcal{C} \wedge D_{\mathcal{B}} \mathcal{C} + \frac{2}{3} \mathcal{C} \wedge \mathcal{C} \wedge \mathcal{C} \right), \quad (3.9)$$

where $D_{\mathcal{B}}$ is the the \mathcal{B} -covariant derivative:

$$D_{\mathcal{B}} \mathcal{C} = d\mathcal{C} + \mathcal{B} \wedge \mathcal{C} + \mathcal{C} \wedge \mathcal{B}. \quad (3.10)$$

This splitting will prove to be a useful guideline in keeping manifest covariance with respect to the diffeomorphism and the non-Abelian gauge transformation.

³Note that if the isometry algebra \mathfrak{g}_i is not associative — as is the case with Poincaré algebra discussed in [15–18] — then the requirement of the closure of the algebra is that the color algebra \mathfrak{g}_c be associative (for the first term in (3.2) to be in the product algebra) and commutative (for the second term in (3.2) to vanish).

⁴In the Introduction, we sketched our model without taking into account this subtraction for the sake of simplicity.

⁵Later, we will take advantage of this decomposition in solving the torsionless condition to convert Chern-Simons formulation into metric formulation.

3.2 Basis of algebra

For further detailed analysis, we set our conventions and notations of the associative algebra involved. The \mathfrak{sl}_2 has three generators J_0, J_1, J_2 . Combining them with the center generator J , one obtains $\mathfrak{gl}_2 = \text{Span}\{J, J_0, J_1, J_2\}$ with the product,

$$J_a J_b = \eta_{ab} J + \epsilon_{abc} J^c \quad [a, b, c = 0, 1, 2]. \quad (3.11)$$

The η_{ab} is the flat metric with mostly positive signs and ϵ_{abc} is the Levi-civita tensor of \mathfrak{sl}_2 with sign convention $\epsilon_{012} = +1$. The generators of the other \mathfrak{gl}_2 will be denoted by \tilde{J}_a and \tilde{J} . In the case of AdS_3 background, the real form of the isometry algebra corresponds to $\mathfrak{so}(2, 2) \simeq \mathfrak{sl}(2, \mathbb{R}) \oplus \mathfrak{sl}(2, \mathbb{R})$, which satisfy

$$(J_a, \tilde{J}_a)^\dagger = -(J_a, \tilde{J}_a), \quad (J, \tilde{J})^\dagger = (J, \tilde{J}). \quad (3.12)$$

In the case of dS_3 background, the real form of the isometry algebra corresponds to $\mathfrak{so}(1, 3) \simeq \mathfrak{sl}(2, \mathbb{C})$, which satisfy

$$(J_a, \tilde{J}_a)^\dagger = -(\tilde{J}_a, J_a), \quad (J, \tilde{J})^\dagger = (\tilde{J}, J). \quad (3.13)$$

Defining the Lorentz generator M_{ab} and the translation generator P_a as

$$M_{ab} = \frac{1}{2} \epsilon_{ab}^c (J_c + \tilde{J}_c), \quad P_a = \frac{1}{2\sqrt{\sigma}} (J_a - \tilde{J}_a), \quad (3.14)$$

where $\sigma = +1$ for AdS_3 and $\sigma = -1$ for dS_3 , we recover the standard commutation relations

$$[M_{ab}, M_{cd}] = 2(\eta_{d[a} M_{b]c} + \eta_{c[b} M_{a]d}), \quad [M_{ab}, P_c] = 2\eta_{c[b} P_{a]}, \quad [P_a, P_b] = \sigma M_{ab}, \quad (3.15)$$

of $\mathfrak{so}(2, 2)$ and $\mathfrak{so}(1, 3)$ for $\sigma = +1$ and -1 , respectively. The reality structure of \mathfrak{gl}_2 determines that of the full algebra \mathfrak{g} in (3.4). As we remarked before, the latter is isomorphic to $\mathfrak{sl}_{2N} \oplus \mathfrak{sl}_{2N}$, hence the conditions (3.12) and (3.13) define which real form of \mathfrak{sl}_{2N} we are dealing with.

The color algebra $\mathfrak{su}(N)$ can be supplemented with the center \mathbf{I} to form the associative algebra $\mathfrak{u}(N)$, with the product

$$\mathbf{T}_I \mathbf{T}_J = \frac{1}{N} \delta_{IJ} \mathbf{I} + (g_{IJ}^K + i f_{IJ}^K) \mathbf{T}_K \quad [I, J, K = 1, \dots, N^2 - 1]. \quad (3.16)$$

The totally symmetric and anti-symmetric structure constants g_{IJK} and f_{IJK} are both real-valued.

We normalize the center generators of both algebras such that their traces are given by⁶

$$\text{Tr}(J) = 2\sqrt{\sigma}, \quad \text{Tr}(\tilde{J}) = -2\sqrt{\sigma}, \quad \text{Tr}(\mathbf{I}) = N. \quad (3.17)$$

The traces of all other elements vanish. This also defines the trace convention in the Chern-Simons action (3.1). With the associative product defined in (3.11), these traces yield all the invariant multilinear forms. For instance, we get the bilinear forms,

$$\text{Tr}(J_a J_b) = 2\sqrt{\sigma} \eta_{ab}, \quad \text{Tr}(\tilde{J}_a \tilde{J}_b) = -2\sqrt{\sigma} \eta_{ab}, \quad \text{Tr}(\mathbf{T}_I \mathbf{T}_J) = \delta_{IJ}, \quad (3.18)$$

which extract the quadratic part of the action.

⁶We use the same notation Tr for the traces of both the isometry algebra and the color algebra.

In the Chern-Simons formulation, the equation of motion is the zero curvature condition: $\mathcal{F} = 0$. In searching for classical solutions, we choose to decompose the subspaces \mathfrak{b} and \mathfrak{c} in (3.5) as

$$\mathfrak{b} = \mathfrak{b}_{\text{GR}} \oplus \mathfrak{b}_{\text{Gauge}}, \quad \mathfrak{c} = \mathfrak{iso} \otimes \mathfrak{su}(N). \quad (3.19)$$

Here, the gravity plus gauge sector corresponds to

$$\mathfrak{b}_{\text{GR}} = \mathfrak{iso} \otimes \mathbf{I}, \quad \mathfrak{b}_{\text{Gauge}} = \text{id} \otimes \mathfrak{su}(N), \quad (3.20)$$

in which \mathfrak{iso} stands for the isometry algebra of the (A)dS₃ space:

$$\mathfrak{iso} = \mathfrak{sl}_2 \oplus \mathfrak{sl}_2. \quad (3.21)$$

There is a trivial vacuum solution where the connection \mathcal{A} is nonzero only for the color-singlet component:

$$\mathcal{B} = \left(\frac{1}{2} \omega^{ab} M_{ab} + \frac{1}{\ell} e^a P_a \right) \mathbf{I}, \quad \mathcal{C} = 0. \quad (3.22)$$

The zero-curvature condition imposes to ω^{ab} and e^a the usual zero (A)dS curvature and zero torsion conditions:

$$d\omega^{ab} + \omega^a{}_c \wedge \omega^{cb} + \frac{\sigma}{\ell^2} e^a \wedge e^b = 0, \quad (3.23)$$

$$de^a + \omega^{ab} \wedge e_b = 0, \quad (3.24)$$

which define the (A)dS₃ space with the radius ℓ , or equivalently with the cosmological constant $\Lambda = -(\sigma/\ell^2)$.

For a general solution, we again decompose $\mathcal{A} = \mathcal{B} + \mathcal{C}$ according to (3.19). The gravity plus gauge sector takes the form

$$\mathcal{B} = \left[\frac{1}{2} \left(\omega^{ab} + \frac{1}{\ell} \Omega^{ab} \right) M_{ab} + \frac{1}{\ell} e^a P_a \right] \mathbf{I} + \mathbf{A} + \tilde{\mathbf{A}}, \quad (3.25)$$

where $\mathbf{A} = A^I J \mathbf{T}_I$ and $\tilde{\mathbf{A}} = \tilde{A}^I \tilde{J} \mathbf{T}_I$ are two copies of $\mathfrak{su}(N)$ gauge field with

$$(\mathbf{A}, \tilde{\mathbf{A}})^\dagger = - \begin{cases} (\mathbf{A}, \tilde{\mathbf{A}}) & [\sigma = +1] \\ (\tilde{\mathbf{A}}, \mathbf{A}) & [\sigma = -1] \end{cases}. \quad (3.26)$$

In (3.25), the splitting $\omega^{ab} + \frac{1}{\ell} \Omega^{ab}$ in the gravity part is arbitrary and is purely for later convenience. The matter sector is composed of

$$\mathcal{C} = \frac{1}{\ell} \left(\varphi^a J_a + \tilde{\varphi}^a \tilde{J}_a \right). \quad (3.27)$$

Here, the colored massless spin-two fields $\varphi^a = \varphi^{a,I} \mathbf{T}_I$ and $\tilde{\varphi}^a = \tilde{\varphi}^{a,I} \mathbf{T}_I$ take value in $\mathfrak{su}(N)$ carrying the adjoint representation. They satisfy

$$(\varphi^a, \tilde{\varphi}^a)^\dagger = + \begin{cases} (\varphi^a, \tilde{\varphi}^a) & [\sigma = +1] \\ (\tilde{\varphi}^a, \varphi^a) & [\sigma = -1] \end{cases}. \quad (3.28)$$

Note that the above has a sign difference from (3.26).

We may find solutions by demanding that (3.25) and (3.27) solve for the zero curvature condition. While this procedure straightforwardly yields nontrivial solutions, for better physical interpretations, we shall first recast the Chern-Simons formulation to the metric formulation and then obtain these nontrivial solutions by solving the latter's field equations.

4 Color-decorated (A)dS₃ gravity: metric formulation

So far, we described the theory in terms of the gauge field \mathcal{A} , so the fact that we are dealing with color-decorated gravity is not tangible. For the sake of concreteness and the advantage of intuitiveness, we shall recast the theory in metric formulation.

We first need to solve the torsionless conditions. This is technically a cumbersome step. Here, we take a short way out from this problem. The idea is that, instead of solving the torsionless conditions for all the colored fields, we shall do it only for the singlet graviton, which we identified above with the metric. This will still allow us to write the action in a metric form but, apart from the gravity, all other colored fields will be still described by a first-order Lagrangian.

In three dimensions, any spectrum with spin greater than zero can be written as a first-order Lagrangian which describes only one helicity mode. If one solves the torsionless conditions for the remaining non-gravity fields, the two fields describing helicity positive and negative modes will combine to generate a single field with a standard second-order Lagrangian. However, this last step appears not necessary and even impossible for certain spectra.

In the following, we will derive the full metric action for the first-order Lagrangian description. For the second-order Lagrangian description, we shall only identify the potential term, leaving aside the explicit form of kinetic terms.

4.1 Colored gravity around singlet vacuum

Starting from the Chern-Simons formulation, described in terms of e^a , $\omega^{ab} + \Omega^{ab}/\ell$, $(\mathbf{A}, \tilde{\mathbf{A}})$ and $(\boldsymbol{\varphi}, \tilde{\boldsymbol{\varphi}})$, we construct a metric formulation by solving the torsionless condition of the gravity sector. This condition is given by

$$de^a + \left(\omega^{ab} + \frac{1}{\ell} \Omega^{ab} \right) \wedge e_b + \frac{\sqrt{\sigma}}{N\ell} \epsilon^{abc} \text{Tr}(\boldsymbol{\varphi}_b \wedge \boldsymbol{\varphi}_c - \tilde{\boldsymbol{\varphi}}_b \wedge \tilde{\boldsymbol{\varphi}}_c) = 0, \quad (4.1)$$

where we require ω^{ab} to satisfy the standard torsionless condition (3.24). This forces $\Omega^{ab} = \epsilon^{abc} \Omega_c$ to satisfy

$$\Omega^{[a} \wedge e^{b]} - \frac{\sqrt{\sigma}}{N} \text{Tr}(\boldsymbol{\varphi}^a \wedge \boldsymbol{\varphi}^b - \tilde{\boldsymbol{\varphi}}^a \wedge \tilde{\boldsymbol{\varphi}}^b) = 0. \quad (4.2)$$

With the above condition together with the standard torsionless condition (3.24), the action (3.1) can be recast to the sum of three parts:

$$S = S_{\text{Gravity}} + S_{\text{CS}} + S_{\text{Matter}}. \quad (4.3)$$

The first term S_{Gravity} is the action for the (A)dS₃ gravity, given by⁷

$$\begin{aligned} S_{\text{Gravity}}[g] &= \frac{\kappa N}{4\pi\ell} \int \epsilon_{abc} e^a \wedge \left(d\omega^{bc} + \omega^{bd} \wedge \omega_d^c + \frac{\sigma}{3\ell^2} e^b \wedge e^c \right) \\ &= \frac{1}{16\pi G} \int d^3x \sqrt{|g|} \left(R + \frac{2\sigma}{\ell^2} \right), \end{aligned} \quad (4.4)$$

where the Chern-Simons level is related to the Newton's constant G , the (A)dS₃ radius ℓ and the rank of the color algebra N by

$$\kappa = \frac{\ell}{4NG}. \quad (4.5)$$

The second term S_{CS} is the Chern-Simons action for $\mathfrak{su}(N) \oplus \mathfrak{su}(N)$ gauge algebra:

$$S_{\text{CS}}[\mathbf{A}\tilde{\mathbf{A}}] = \frac{\kappa\sqrt{\sigma}}{2\pi} \int \left[\text{Tr} \left(\mathbf{A} \wedge d\mathbf{A} + \frac{2}{3} \mathbf{A} \wedge \mathbf{A} \wedge \mathbf{A} \right) - \text{Tr} \left(\tilde{\mathbf{A}} \wedge d\tilde{\mathbf{A}} + \frac{2}{3} \tilde{\mathbf{A}} \wedge \tilde{\mathbf{A}} \wedge \tilde{\mathbf{A}} \right) \right]. \quad (4.6)$$

In the uncolored Chern-Simons gravity, it is unclear whether the Chern-Simons level κ has to be quantized since the gauge group is not compact. However, in the case of colored Chern-Simons gravity, the level κ should take an integer value for the consistency of S_{CS} (4.6) under a large $\text{SU}(N) \times \text{SU}(N)$ gauge transformation.

The last term S_{Matter} is the action for the colored massless spin-two fields φ^a and $\tilde{\varphi}^a$. To derive it, we use the decompositions (3.25) and (3.27), and simplify by using (4.2). We get

$$S_{\text{Matter}}[\varphi, \tilde{\varphi}] = \frac{1}{16\pi G} \int \left[\frac{1}{N} L[\varphi, \tilde{\varphi}, \ell] - \frac{1}{\ell^2} \epsilon_{abc} e^a \wedge \Omega^b(\varphi, \tilde{\varphi}) \wedge \Omega^c(\varphi, \tilde{\varphi}) \right], \quad (4.7)$$

where the three-form Lagrangian $L[\varphi, \tilde{\varphi}; \ell]$ is given by

$$\begin{aligned} L[\varphi, \tilde{\varphi}, \ell] &= L_+[\varphi, \ell] - L_-[\tilde{\varphi}, \ell], \\ L_{\pm}[\varphi, \ell] &= 2\sqrt{\sigma} \text{Tr} \left[\frac{1}{\ell} \varphi_a \wedge D\varphi^a + \frac{1}{\ell^2} \epsilon_{abc} \left(\frac{\pm 1}{\sqrt{\sigma}} e^a \wedge \varphi^b \wedge \varphi^c + \frac{2}{3} \varphi^a \wedge \varphi^b \wedge \varphi^c \right) \right]. \end{aligned} \quad (4.8)$$

In this expression, the covariant derivative D is with respect to both the Lorentz transformation and the $\mathfrak{su}(N)$ gauge transformation:

$$\begin{aligned} D\varphi^a &= d\varphi^a + \omega^{ab} \wedge \varphi_b + \mathbf{A} \wedge \varphi^a + \varphi^a \wedge \mathbf{A}, \\ D\tilde{\varphi}^a &= d\tilde{\varphi}^a + \omega^{ab} \wedge \tilde{\varphi}_b + \tilde{\mathbf{A}} \wedge \tilde{\varphi}^a + \tilde{\varphi}^a \wedge \tilde{\mathbf{A}}. \end{aligned} \quad (4.9)$$

The last term in (4.7) is an implicit function of φ^a and $\tilde{\varphi}^a$. It is proportional to

$$\epsilon_{abc} e^a \wedge \Omega^b \wedge \Omega^c = \frac{1}{3} \epsilon_{abc} e^a \wedge e^b \wedge e^c \Omega_{[d}{}^d \Omega_{e]}{}^e, \quad (4.10)$$

where $\Omega^a = \Omega_b{}^a e^b$. From (4.2), they are determined to be

$$\Omega_a{}^b = \frac{1}{N} W_a{}^b(\varphi, \tilde{\varphi}) = \Omega_a{}^b(\varphi, \tilde{\varphi}), \quad (4.11)$$

⁷In our normalization, $d^3x\sqrt{|g|} = \frac{1}{6} \epsilon_{abc} e^a \wedge e^b \wedge e^c$.

where $W_a{}^b(\varphi, \tilde{\varphi})$ is given by

$$\begin{aligned} W_a{}^b(\varphi, \tilde{\varphi}) &= W_a{}^b(\varphi) - W_a{}^b(\tilde{\varphi}), \\ W_a{}^b(\varphi) &= 4\sqrt{\sigma} \text{Tr} \left(\varphi_{[a}{}^b \varphi_{c]}{}^c - \frac{1}{4} \delta_a^b \varphi_{[c}{}^c \varphi_{d]}{}^d \right). \end{aligned} \quad (4.12)$$

Here, $\varphi_b{}^a$ are the components of φ^a : $\varphi^a = \varphi_b{}^a e^b$. Notice that only the term (4.10) — which is quartic in φ^a and $\tilde{\varphi}^a$ — gives the cross couplings between φ 's and $\tilde{\varphi}$'s.

4.2 First-order description

Gathering all above results and replacing the dreibein e^a in terms of the metric $g_{\mu\nu}$, the colored gravity action reads

$$S = S_{\text{CS}} + \frac{1}{16\pi G} \int d^3x \sqrt{|g|} \left[R - V(\varphi, \tilde{\varphi}) + \frac{2\sqrt{\sigma}}{N\ell} \epsilon^{\mu\nu\rho} \text{Tr} \left(\varphi_\mu{}^\lambda D_\nu \varphi_{\rho\lambda} - \tilde{\varphi}_\mu{}^\lambda D_\nu \tilde{\varphi}_{\rho\lambda} \right) \right], \quad (4.13)$$

where the covariant derivative is given by

$$D_\mu \varphi_{\nu\rho} = \nabla_\mu \varphi_{\nu\rho} + [\mathbf{A}_\mu, \varphi_{\nu\rho}] \quad (4.14)$$

and the scalar potential function is given by

$$\begin{aligned} V(\varphi, \tilde{\varphi}) &= -\frac{1}{N\ell^2} \text{Tr} \left[2\sigma \mathbf{I} + 4 \left(\varphi_{[\mu}{}^\mu \varphi_{\nu]}{}^\nu + \tilde{\varphi}_{[\mu}{}^\mu \tilde{\varphi}_{\nu]}{}^\nu \right) + 8\sqrt{\sigma} \left(\varphi_{[\mu}{}^\mu \varphi_{\nu]}{}^\nu \varphi_{\rho]}{}^\rho - \tilde{\varphi}_{[\mu}{}^\mu \tilde{\varphi}_{\nu]}{}^\nu \tilde{\varphi}_{\rho]}{}^\rho \right) \right] \\ &\quad - \frac{16\sigma}{N^2\ell^2} \text{Tr} \left(\varphi_{[\mu}{}^\nu \varphi_{\rho]}{}^\rho - \tilde{\varphi}_{[\mu}{}^\nu \tilde{\varphi}_{\rho]}{}^\rho \right) \text{Tr} \left(\varphi_{[\nu}{}^\mu \varphi_{\lambda]}{}^\lambda - \tilde{\varphi}_{[\nu}{}^\mu \tilde{\varphi}_{\lambda]}{}^\lambda \right) \\ &\quad + \frac{6\sigma}{N^2\ell^2} \left[\text{Tr} \left(\varphi_{[\mu}{}^\mu \varphi_{\nu]}{}^\nu - \tilde{\varphi}_{[\mu}{}^\mu \tilde{\varphi}_{\nu]}{}^\nu \right) \right]^2. \end{aligned} \quad (4.15)$$

The scalar potential function consists of single-trace and double-trace parts. The single-trace part originates from the Chern-Simons cubic interaction, while the double-trace part originates from solving the torsionless conditions. For a general configuration, all terms in the potential have the same order in large N as the other terms in (4.13).

Already at this stage, the content of the colored gravity is clearly demonstrated: it is a theory of colored massless left-moving and right-moving spin-two fields, as seen from the kinetic term in (4.8) or (4.13). They interact covariantly with the color singlet gravity and also with the Chern-Simons color gauge fields. Moreover, they interact with each other through the potential $V(\varphi, \tilde{\varphi})$. The self-interaction is governed by the constant $1/N$. The single-trace cubic interaction is stronger than the gravitational cubic interaction by the factor of \sqrt{N} . Therefore, at large N and for fixed Newton's constant, the colored massless spin-two fields will be strongly coupled to each other.

4.3 Second-order description

In principle, we could also solve the torsionless condition for the colored spin-two fields and obtain a second-order Lagrangian (although this spoils the minimal interactions to the $\mathfrak{su}(N)$ gauge fields \mathbf{A} and $\tilde{\mathbf{A}}$). It amounts to taking linear combinations

$$\chi_{\mu\nu} = \sqrt{\sigma} (\varphi_{\mu\nu} - \tilde{\varphi}_{\mu\nu}), \quad \tau_{\mu\nu} = \varphi_{\mu\nu} + \tilde{\varphi}_{\mu\nu}, \quad (4.16)$$

and integrating out the *torsion* part $\tau_{\mu\nu}$, while keeping $\chi_{\mu\nu}$. The resulting action is given by

$$S = S_{\text{CS}} + \frac{1}{16\pi G} \int d^3x \sqrt{|g|} \left[R - V(\chi) + \mathcal{L}_{\text{CM}}(\chi, \nabla\chi, \mathbf{A}, \tilde{\mathbf{A}}) \right]. \quad (4.17)$$

The Lagrangian \mathcal{L}_{CM} reads

$$\mathcal{L}_{\text{CM}}(\chi, \nabla\chi, \mathbf{A}, \tilde{\mathbf{A}}) = \frac{1}{N} \text{Tr} \left(2 \chi_{\mu\nu} \nabla^2 \chi^{\mu\nu} + \dots \right), \quad (4.18)$$

where the ellipses include other tensor contractions together with higher-order terms of the form, $\chi^n (\nabla\chi)^2$ with $n \geq 1$ as well as couplings to the gauge fields \mathbf{A} and $\tilde{\mathbf{A}}$. We do not attempt to obtain the complete structure of these derivative terms.

The potential function $V(\chi)$ corresponds to the extremum of

$$\begin{aligned} V(\chi, \tau) = & -\frac{2\sigma}{N\ell^2} \text{Tr} \left(\mathbf{I} + \chi_{[\mu}{}^\mu \chi_{\nu]}{}^\nu + \sigma \tau_{[\mu}{}^\mu \tau_{\nu]}{}^\nu + \chi_{[\mu}{}^\mu \chi_{\nu}{}^\nu \chi_{\rho]}{}^\rho + 3\sigma \chi_{[\mu}{}^\mu \tau_{\nu}{}^\nu \tau_{\rho]}{}^\rho \right) \\ & - \frac{4}{N^2 \ell^2} \text{Tr} \left(\chi_{[\mu}{}^\nu \tau_{\rho]}{}^\rho + \tau_{[\mu}{}^\nu \chi_{\rho]}{}^\rho \right) \text{Tr} \left(\chi_{[\nu}{}^\mu \tau_{\lambda]}{}^\lambda + \tau_{[\nu}{}^\mu \chi_{\lambda]}{}^\lambda \right) \\ & + \frac{6}{N^2 \ell^2} \left[\text{Tr} \left(\chi_{[\mu}{}^\mu \tau_{\nu]}{}^\nu \right) \right]^2, \end{aligned} \quad (4.19)$$

along the $\tau_{\mu\nu}$ direction. As the extremum equation for $\tau_{\mu\nu}$ is linear in $\tau_{\mu\nu}$,

$$\mathcal{M}(\chi) \cdot \tau = 0, \quad (4.20)$$

it must be that the unique solution is $\tau_{\mu\nu} = 0$ for a generic configuration of $\chi_{\mu\nu}$.⁸ Proceeding with this situation, we end up with the cubic potential for $\chi_{\mu\nu}$:

$$V(\chi) = -\frac{2\sigma}{N\ell^2} \text{Tr} \left(\mathbf{I} + \chi_{[\mu}{}^\mu \chi_{\nu]}{}^\nu + \chi_{[\mu}{}^\mu \chi_{\nu}{}^\nu \chi_{\rho]}{}^\rho \right). \quad (4.21)$$

This potential has a noticeably simple form, but also has rich implications as we shall discuss in the next sections.

5 Classical vacua of colored gravity

5.1 Identification of vacuum solutions

Having identified the action in metric formulation, we now search for classical vacua that solve the field equations of motion:

$$-\frac{\delta \mathcal{L}_{\text{CM}}}{\delta g_{\mu\nu}} = G_{\mu\nu} - \frac{1}{2} V(\chi) g_{\mu\nu}, \quad \frac{\delta \mathcal{L}_{\text{CM}}}{\delta \chi_{\mu\nu}} = \frac{\partial V(\chi)}{\partial \chi_{\mu\nu}}, \quad (5.1)$$

$$-\frac{N}{2\sqrt{\sigma}\ell} \frac{\delta \mathcal{L}_{\text{CM}}}{\delta \mathbf{A}_\mu} = \epsilon^{\mu\nu\rho} \mathbf{F}_{\nu\rho}, \quad -\frac{N}{2\sqrt{\sigma}\ell} \frac{\delta \mathcal{L}_{\text{CM}}}{\delta \tilde{\mathbf{A}}_\mu} = \epsilon^{\mu\nu\rho} \tilde{\mathbf{F}}_{\nu\rho}. \quad (5.2)$$

In order to find their solutions, we assume that the colored massless spin-two fields are covariantly constant with the trivial $\mathfrak{su}(N)$ gauge connection,

$$\mathbf{A} = 0, \quad \tilde{\mathbf{A}} = 0, \quad \nabla_\rho \chi_{\mu\nu} = 0. \quad (5.3)$$

⁸There can also exist nontrivial $\tau_{\mu\nu}$ solutions at special values of $\chi_{\mu\nu}$, corresponding to kernel of \mathcal{M} in (4.20). They break the parity symmetry spontaneously, and hence of special interest. We relegate complete classification of these null solutions in a separate paper.

This can be satisfied by

$$\chi_{\mu\nu} = g_{\mu\nu} \mathbf{X} \quad \text{for} \quad \mathbf{X} = \text{constant} \in \mathfrak{su}(N). \quad (5.4)$$

Physically, we interpret this as the colored spin-two matter acting as Higgs field. In Poincaré invariant field theory, the vacuum is Poincaré invariant, so only a scalar field φ (which is proportional to an identity operator $\varphi \propto \mathbb{I}$) can take a vacuum expectation value, $\langle \varphi \rangle = v$. On the other hand, fields with nonzero spin cannot develop a nonzero expectation value since it is incompatible with the Lorentz invariance. In generally covariant field theory, where the background metric $g_{\mu\nu}$ plays the role of first fundamental form, the spin-2 field $\chi_{\mu\nu}$ can similarly develop a nonzero vacuum expectation value $\langle \chi_{\mu\nu} \rangle = v g_{\mu\nu}$ proportional to the metric $g_{\mu\nu}$, while all other fields of higher spin cannot. We thus refer this phenomenon to as ‘gravitational Higgs mechanism’.

With (5.3) and (5.4), the equations in the second line (5.2) trivialize and the rest reduce to

$$G_{\mu\nu} - \frac{1}{2} V(\mathbf{X}) g_{\mu\nu} = 0 \quad \text{and} \quad \frac{\partial V(\mathbf{X})}{\partial \mathbf{X}} = 0, \quad (5.5)$$

where $V(\mathbf{X}) = V(\chi_{\mu\nu} = g_{\mu\nu} \mathbf{X})$ is given by

$$V(\mathbf{X}) = -\frac{2\sigma}{N\ell^2} \text{Tr}(\mathbf{I} + 3\mathbf{X}^2 + \mathbf{X}^3). \quad (5.6)$$

From (5.5), the extremum of the potential defines the corresponding cosmological constant:

$$\Lambda = \frac{1}{2} V(\mathbf{X}). \quad (5.7)$$

Although cubic, being a matrix-valued function, the potential $V(\mathbf{X})$ may admit a large number of nontrivial extrema that depends on the color algebra $\mathfrak{su}(N)$. If exists, each of such extrema will define a distinct vacuum with a different cosmological constant (5.7). As an illustration of this potential, consider the function $f(\mathbf{X}) = \frac{1}{N} \text{Tr}(\mathbf{I} + 3\mathbf{X}^2 + \mathbf{X}^3)$ for the \mathbf{X} belonging to $\mathfrak{su}(3)$. The 3×3 matrix \mathbf{X} can be diagonalized by a $\text{SU}(3)$ rotation to

$$\mathbf{X} = a \begin{bmatrix} 1 & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & -2 \end{bmatrix} + b \begin{bmatrix} -2 & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & 1 \end{bmatrix}. \quad (5.8)$$

We plot the function $f(a, b)$ in figure 1. It clearly exhibits four extremum points: $(0, 0)$, $(2, 0)$, $(0, 2)$ and $(-2, -2)$. The first point at the origin gives $f = 1$, whereas the other three points all give $f = 9$. In fact, these three points are connected by $\text{SU}(3)$ rotation.

We now explicitly identify the extrema of potential function (5.6) for arbitrary value of N . The extremum points are defined by the equation:

$$\delta V(\mathbf{X}) = -\frac{6\sigma}{N\ell^2} \text{Tr}[(2\mathbf{X} + \mathbf{X}^2)\delta\mathbf{X}] = 0. \quad (5.9)$$

Since \mathbf{X} is traceless, it follows that $\delta\mathbf{X}$ is also traceless. Thus, the equation reads

$$2\mathbf{X} + \mathbf{X}^2 = \frac{1}{N} \text{Tr}(2\mathbf{X} + \mathbf{X}^2) \mathbf{I}. \quad (5.10)$$

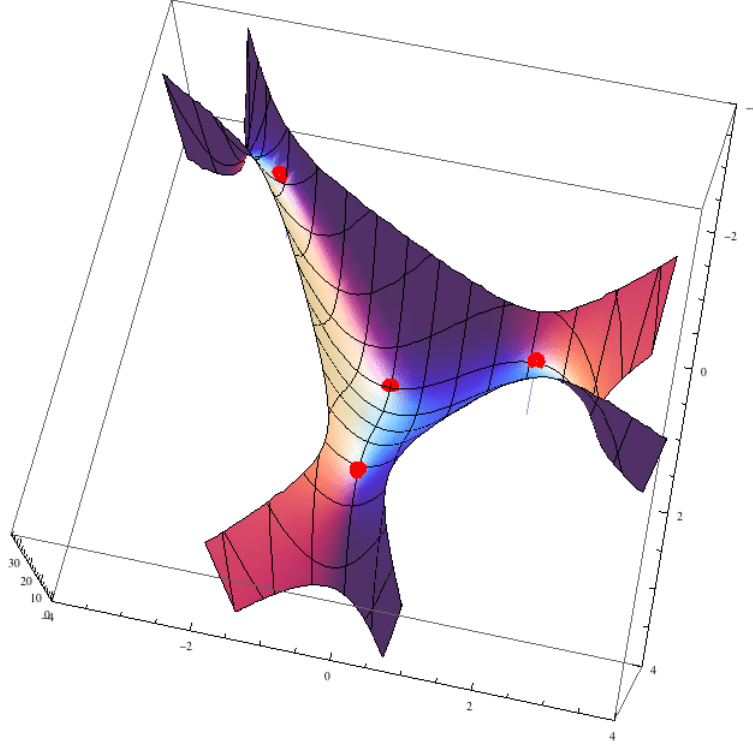


Figure 1. The shape of the potential function for $\mathfrak{su}(3)$.

Since $\text{Tr}(\mathbf{I} + \mathbf{X})^2 \neq 0$ — otherwise it would follow from (5.10) that the matrix $\mathbf{I} + \mathbf{X}$ is nilpotent while having a non-trivial trace — one can redefine the matrix \mathbf{X} in terms of \mathbf{Z} :

$$\mathbf{Z} = \sqrt{\frac{N}{\text{Tr}(\mathbf{I} + \mathbf{X})^2}} (\mathbf{I} + \mathbf{X}), \quad (5.11)$$

or equivalently,

$$\mathbf{X} = \frac{N}{\text{Tr}(\mathbf{Z})} \mathbf{Z} - \mathbf{I}. \quad (5.12)$$

This simplifies the equation (5.10) as

$$\mathbf{Z}^2 = \mathbf{I}. \quad (5.13)$$

Complete solutions of this equation, up to $\text{SU}(N)$ rotations, are given by

$$\mathbf{Z}_k = \begin{bmatrix} \mathbf{I}_{(N-k) \times (N-k)} & 0 \\ 0 & -\mathbf{I}_{k \times k} \end{bmatrix}, \quad k = 0, 1, \dots, \left\lfloor \frac{N-1}{2} \right\rfloor. \quad (5.14)$$

where the upper bound of k is fixed by $\left\lfloor \frac{N-1}{2} \right\rfloor$ due to the property that \mathbf{X}_{N-k} is a $\text{SU}(N)$ rotation of \mathbf{X}_k . Notice also that, when N is even, $k = \frac{N}{2}$ is excluded since it leads to $\text{Tr}(\mathbf{Z}) = 0$ for which \mathbf{X} is ill-defined. Plugging the solutions (5.14) to the potential, we can identify the values of the potential at the extrema as

$$V(\mathbf{X}_k) = -\frac{2\sigma}{\ell^2} \left(\frac{N}{\text{Tr}(\mathbf{Z}_k)} \right)^2 = -\frac{2\sigma}{\ell^2} \left(\frac{N}{N-2k} \right)^2. \quad (5.15)$$

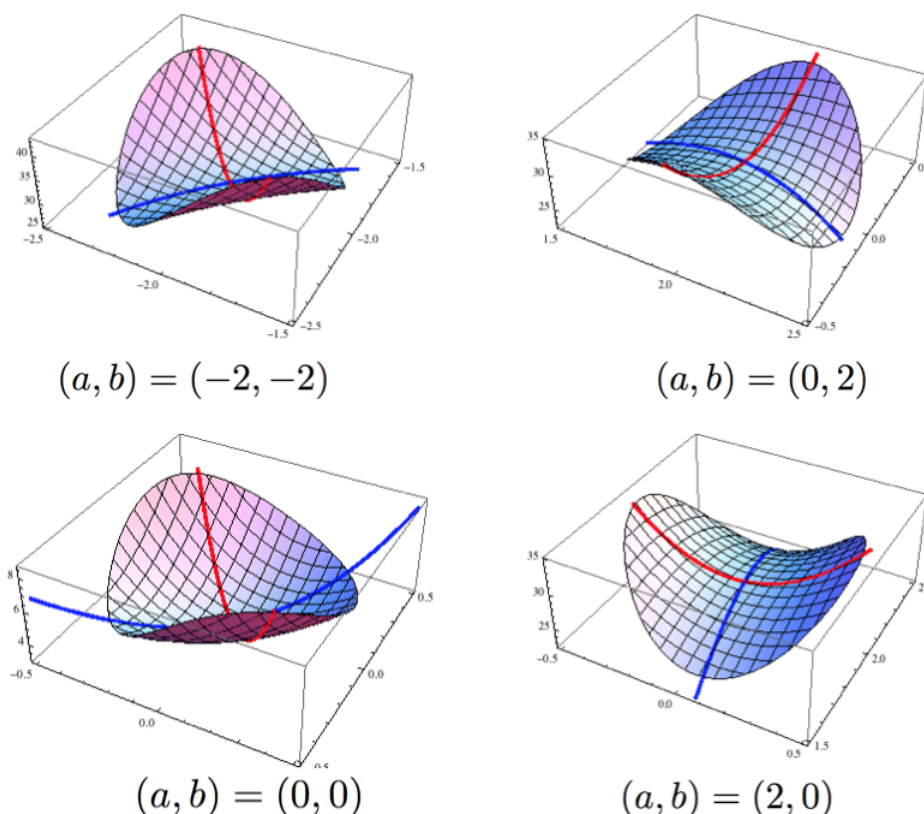


Figure 2. Potentials around each rainbow vacua with $N = 3$. The $(a, b) = (0, 0)$ vacuum is the minimum/maximum in dS_3/AdS_3 space. The others, connected by $SU(3)$ transformation, are all saddle points.

These values play the role of the cosmological constant at the k -th extremum, according to (5.7).

Let us discuss more on the potential (5.6). Firstly, the cubic form shows that the potential is not bounded from below or above. Secondly, the overall factor σ shows that the overall sign of the potential depends whether we consider AdS_3 or dS_3 background. Thirdly, we can understand better the stability of the extrema we found by considering the second variation of the potential,

$$\delta^2 V(\mathbf{X}_k) = -\frac{12\sigma}{(N-2k)\ell^2} \text{Tr}(\mathbf{Z}_k \delta \mathbf{X}^2). \quad (5.16)$$

The Hessian is not positive(or negative)-definite for an arbitrary $\delta \mathbf{X}$ except the singlet vacuum $k = 0$. So, all $k \neq 0$ vacua are saddle points and the $k = 0$ vacuum is the minimum/maximum in dS_3/AdS_3 space.

5.2 $N = 3$ example and linearized spectrum

In the standard Higgs mechanism, the gauge fields combine with the Goldstone bosons to become massive vector fields. In the following, we will analyze the analogous mechanism in our model of colored gravity. For the concreteness, let us consider the $k = 1$ vacuum solution (5.14) in $N = 3$ case (see figure 2). This solution has a non-zero background for

the colored matter fields which breaks the SU(3) symmetry down to SU(2) \times U(1). We linearize the colored matter fields $(\varphi, \tilde{\varphi})$ around this vacuum as

$$\varphi_{\mu\nu} = \frac{\mathbf{X}_1}{2\sqrt{\sigma}} g_{\mu\nu} + \varphi_{\mu\nu}^{\text{fluc}}, \quad \tilde{\varphi}_{\mu\nu} = -\frac{\mathbf{X}_1}{2\sqrt{\sigma}} g_{\mu\nu} + \tilde{\varphi}_{\mu\nu}^{\text{fluc}}, \quad (5.17)$$

where the background value of $(\varphi, \tilde{\varphi})$ is proportional to the matrix \mathbf{X}_1 (5.12), whose explicit form reads

$$\mathbf{X}_1 = \begin{bmatrix} 2 & 0 & 0 \\ 0 & 2 & 0 \\ 0 & 0 & -4 \end{bmatrix}. \quad (5.18)$$

The fluctuation parts of the colored matter fields and the spin-one Cherns-Simons gauge fields can be decomposed as

$$\begin{aligned} \varphi_{\mu\nu}^{\text{fluc}} &= \frac{3}{\sqrt{2}} \varrho_{\mu\nu}^a \mathbf{T}_{\text{su}(2)}^a + \sqrt{\frac{3}{2}} (\tilde{\psi}_{\mu\nu} - 2\psi_{\mu\nu}) \mathbf{T}_{\text{u}(1)} + \frac{3}{\sqrt{2}} \phi_{\mu\nu}^i \mathbf{T}_{\text{BS}}^i, \\ \tilde{\varphi}_{\mu\nu}^{\text{fluc}} &= \frac{3}{\sqrt{2}} \tilde{\varrho}_{\mu\nu}^a \mathbf{T}_{\text{su}(2)}^a + \sqrt{\frac{3}{2}} (\psi_{\mu\nu} - 2\tilde{\psi}_{\mu\nu}) \mathbf{T}_{\text{u}(1)} + \frac{3}{\sqrt{2}} \tilde{\phi}_{\mu\nu}^i \mathbf{T}_{\text{BS}}^i, \end{aligned} \quad (5.19)$$

$$\begin{aligned} \mathbf{A}_\mu &= A_\mu^a \mathbf{T}_{\text{su}(2)}^a + A_\mu \mathbf{T}_{\text{u}(1)} + \frac{1}{\sqrt{8}} A_\mu^i \mathbf{Z}_1 \mathbf{T}_{\text{BS}}^i, \\ \tilde{\mathbf{A}}_\mu &= \tilde{A}_\mu^a \mathbf{T}_{\text{su}(2)}^a + \tilde{A}_\mu \mathbf{T}_{\text{u}(1)} + \frac{1}{\sqrt{8}} \tilde{A}_\mu^i \mathbf{Z}_1 \mathbf{T}_{\text{BS}}^i, \end{aligned} \quad (5.20)$$

in terms of the SU(3) generators:

$$\begin{aligned} \mathbf{T}_{\text{su}(2)}^a &= \frac{1}{\sqrt{2}} \begin{bmatrix} \sigma^a & 0 \\ 0 & 0 \end{bmatrix}, \quad \mathbf{T}_{\text{u}(1)} = \frac{1}{\sqrt{6}} \begin{bmatrix} 1 & 0 & 0 \\ 0 & 1 & 0 \\ 0 & 0 & -2 \end{bmatrix}, \\ \mathbf{T}_{\text{BS}}^1 &= \frac{1}{\sqrt{2}} \begin{bmatrix} 0 & 0 & 1 \\ 0 & 0 & 0 \\ 1 & 0 & 0 \end{bmatrix}, \quad \mathbf{T}_{\text{BS}}^2 = \frac{1}{\sqrt{2}} \begin{bmatrix} 0 & 0 & -i \\ 0 & 0 & 0 \\ i & 0 & 0 \end{bmatrix}, \quad \mathbf{T}_{\text{BS}}^3 = \frac{1}{\sqrt{2}} \begin{bmatrix} 0 & 0 & 0 \\ 0 & 0 & 1 \\ 0 & 1 & 0 \end{bmatrix}, \quad \mathbf{T}_{\text{BS}}^4 = \frac{1}{\sqrt{2}} \begin{bmatrix} 0 & 0 & 0 \\ 0 & 0 & -i \\ 0 & i & 0 \end{bmatrix}, \end{aligned} \quad (5.21)$$

where $\sigma^1, \sigma^2, \sigma^3$ are the Pauli matrices. Various factors in (5.19) and (5.20) have been introduced for latter convenience. By plugging (5.19) into the original action (4.13) and expanding the action up to quadratic order in the fluctuations, we obtain the perturbative Lagrangian around the $k=1$ vacuum. We first expand the potential as

$$V(\varphi, \tilde{\varphi}) = -\frac{2}{\ell_1^2} \left(\sigma + \varrho_{[\mu}^a \varrho_{\nu]}^a + \tilde{\varrho}_{[\mu}^a \tilde{\varrho}_{\nu]}^a - \psi_{[\mu} \psi_{\nu]} - \tilde{\psi}_{[\mu} \tilde{\psi}_{\nu]} \right) + \mathcal{O}(\Phi^3), \quad (5.22)$$

and the kinetic part as

$$\begin{aligned} &\frac{2\sqrt{\sigma}}{9\ell_1} \epsilon^{\mu\nu\rho} \text{Tr} \left(\varphi_\mu^\lambda D_\nu \varphi_{\rho\lambda} - \tilde{\varphi}_\mu^\lambda D_\nu \tilde{\varphi}_{\rho\lambda} \right) \\ &= \frac{\sqrt{\sigma}}{\ell_1} \epsilon^{\mu\nu\rho} \left(\varphi_\mu'^{a\lambda} \nabla_\nu \varphi_{\rho\lambda}'^a + \psi_\mu^\lambda \nabla_\nu \psi_{\rho\lambda} + \phi_\mu^i \nabla_\nu \phi_{\rho\lambda}^i + \frac{1}{\sqrt{\sigma}} \phi_{\mu\nu}^i A_\rho^i \right) + c.c + \mathcal{O}(\Phi^3). \end{aligned} \quad (5.23)$$

Here the $\ell_1 = \ell/3$ is the radius of the $k = 1$ (A)dS solution, and $\mathcal{O}(\Phi^3)$ means the cubic-order terms in the fluctuation fields. Combining (5.22) and (5.23), the colored gravity action (4.13) becomes

$$S = S_{\text{CS}} + \frac{1}{16\pi G} \int d^3 \sqrt{|g|} \left(R + \frac{2\sigma}{\ell_1^2} + \mathcal{L}_{\text{RS}} + \mathcal{L}_{\text{BS}} + \mathcal{O}(\Phi^3) \right), \quad (5.24)$$

where the Lagrangian for the residual symmetry part is given by

$$\begin{aligned} \ell_1^2 \mathcal{L}_{\text{RS}} = & \sqrt{\sigma} \ell_1 \epsilon^{\mu\nu\rho} \varrho_\mu^{a\lambda} \nabla_\nu \varrho_{\rho\lambda}^a + 2 \varrho_{[\mu}^{a\ \nu} \varrho_{\nu]}^{a\ \rho} + \text{c.c.} \\ & + \sqrt{\sigma} \ell_1 \epsilon^{\mu\nu\rho} \psi_\mu^\lambda \nabla_\nu \psi_{\rho\lambda} - 2 \psi_{[\mu}^\lambda \psi_{\nu]}^\rho + \text{c.c.} \end{aligned} \quad (5.25)$$

and that for the broken symmetry part by

$$\ell_1^2 \mathcal{L}_{\text{BS}} = \sqrt{\sigma} \ell_1 \epsilon^{\mu\nu\rho} \left(\phi_\mu^{i\lambda} \nabla_\nu \phi_{\rho\lambda}^i + \frac{1}{\sqrt{\sigma}} \phi_{\mu\nu}^i A_\rho^i \right) + \text{c.c.} \quad (5.26)$$

Several remarks are in order:

- In the Lagrangian \mathcal{L}_{RS} (5.25), the fields $\varrho_{\mu\nu}^a$ — associated with the $\mathfrak{su}(2)$ generators — describe the standard massless spin-two fields. On the contrary, the field $\psi_{\mu\nu}$ — associated with the $\mathfrak{u}(1)$ generator — describes a *ghost* massless spin-two due to the sign flip of the no-derivative term (why this sign determines whether the spectrum is ghost or not is explained in appendix A).
- In the Lagrangian \mathcal{L}_{BS} (5.26) — associated with the broken part of the symmetry — has an unusual cross term with the spin-one Cherns-Simons gauge field A_μ^i . In fact, A_μ^i behaves as a Stueckelberg field hence can be removed by a spin-two gauge transformation. Let us remark that this gauge choice is analogous to the unitary gauge in the standard Higgs mechanism. As a result, the Chern-Simons action S_{CS} reduces from $\text{SU}(3)$ to $\text{SU}(2) \times \text{U}(1)$, and the field $\phi_{\mu\nu}^i$ inherits the gauge symmetries of A_μ^i as a second-derivative form:

$$\delta \phi_{\mu\nu}^i = \left(\nabla_\mu \nabla_\nu - \frac{\sigma}{\ell_1^2} g_{\mu\nu} \right) \xi^i. \quad (5.27)$$

This spectrum clearly combines the massless spin-two mode with the spin-one mode in an irreducible manner. It actually corresponds to so-called partially-massless spin-two field [31, 32]. Since our system is after all a Chern-Simons theory, there is no propagating DoF such as a scalar field. Hence, it is clear that we cannot have a massive spin-two as a result of symmetry breaking because it would require not only spin-one but also a scalar mode. We postpone more detailed analysis to the next section.

6 Colored gravity around rainbow vacua

We learned that there are $\lfloor \frac{N+1}{2} \rfloor$ many distinct vacua having different cosmological constants. In this section, we study the colored gravity around each of these vacua and analyze the spectrum. In principle, we can proceed in the same way as we did for the $N = 3$ example in section 5.2, but there is a more systematic way relying on the Chern-Simons formulation.

6.1 Decomposition of algebra revisited

For an efficient treatment of the colored gravity at each distinct vacuum in the Chern-Simons formulation, it is important to identify the proper decomposition of the algebra (3.5). For that, we revisit the isometry and the color algebra decompositions. The isometry algebra can be divided into the rotation part \mathcal{M} and the translation part \mathcal{P} as

$$\mathfrak{iso} = \mathcal{M} \oplus \mathcal{P}, \quad (6.1)$$

the same as the trivial vacuum. For the color algebra, each vacuum spontaneously breaks the Chan-Paton $\mathfrak{su}(N)$ gauge symmetry down to $\mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1)$, and hence the original algebra admits the decomposition:

$$\mathfrak{su}(N) \simeq \mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1) \oplus \mathfrak{bs}. \quad (6.2)$$

Here, \mathfrak{bs} is the vector space corresponding to the *broken symmetry*, spanned by $2k(N-k)$ generators. It is important to note that each part commutes or anti-commutes with the background matrix \mathbf{Z}_k (5.14) as

$$[\mathbf{Z}_k, \mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1)] = 0, \quad \{\mathbf{Z}_k, \mathfrak{bs}\} = 0. \quad (6.3)$$

We now decompose the entire algebra (3.4) according to (3.5) in terms of the gravity plus gauge sector \mathfrak{b} and the matter sector \mathfrak{c} . The former has again two parts similarly to the singlet vacuum case as $\mathfrak{b} = \mathfrak{b}_{\text{GR}} \oplus \mathfrak{b}_{\text{Gauge}}$, but the algebras to which the gravity and the gauge sectors correspond differ from (3.20). They are

$$\mathfrak{b}_{\text{GR}} = (\mathcal{M} \otimes \mathbf{I}) \oplus (\mathcal{P} \otimes \mathbf{Z}_k), \quad \mathfrak{b}_{\text{Gauge}} = \text{id} \otimes (\mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1)). \quad (6.4)$$

The gauge sector is concerned only with the unbroken part of the color algebra. The algebra of the gravity sector is deformed by \mathbf{Z}_k , but still satisfies the same commutation relations with the generators:

$$\mathbf{M}_{ab} = M_{ab} \mathbf{I}, \quad \mathbf{P}_a = P_a \mathbf{Z}_k. \quad (6.5)$$

The one-form gauge fields associated with these sectors are given correspondingly by

$$\begin{aligned} \mathcal{B}_{\text{GR}} &= \frac{1}{2} (\omega^{ab} + \Omega^{ab}) M_{ab} + \frac{1}{\ell_k} e^a P_a, \\ \mathcal{B}_{\text{Gauge}} &= \mathbf{A}_+ + \mathbf{A}_- + \tilde{\mathbf{A}}_+ + \tilde{\mathbf{A}}_- + (\mathbf{A} + \tilde{\mathbf{A}}) \mathbf{Y}_k, \end{aligned} \quad (6.6)$$

where the k -vacuum radius ℓ_k is related to the singlet one as

$$\ell_k := \left(\frac{N-2k}{N} \right) \ell, \quad (6.7)$$

and \mathbf{Y}_k is the traceless matrix:

$$\mathbf{Y}_k = \frac{k \mathbf{I}_+ - (N-k) \mathbf{I}_-}{N}. \quad (6.8)$$

Here again, the spin connection ω^{ab} is the standard one satisfying (3.24), whereas Ω^{ab} will be determined in terms of other fields from the torsionless conditions. The gauge fields \mathbf{A}_\pm and $\tilde{\mathbf{A}}_\pm$ take values in $\mathfrak{su}(N-k)$ for the subscript $+$ and $\mathfrak{su}(k)$ for the subscript $-$, whereas A and \tilde{A} are Abelian gauge fields taking values in $\mathfrak{u}(1)$.

In the case of non-singlet vacua, the matter sector space has two parts:

$$\mathfrak{c} = \mathfrak{c}_{\text{CM}} \oplus \mathfrak{c}_{\text{BS}}. \quad (6.9)$$

For the introduction of each elements, let us first define the generators of $\mathfrak{gl}_2 \oplus \mathfrak{gl}_2$ deformed by \mathbf{Z}_k as

$$\begin{aligned} \mathbf{J}_a &= J_a \mathbf{I}_+ + \tilde{J}_a \mathbf{I}_-, & \mathbf{J} &= J \mathbf{I}_+ + \tilde{J} \mathbf{I}_-, \\ \tilde{\mathbf{J}}_a &= J_a \mathbf{I}_- + \tilde{J}_a \mathbf{I}_+, & \tilde{\mathbf{J}} &= J \mathbf{I}_- + \tilde{J} \mathbf{I}_+, \end{aligned} \quad (6.10)$$

where \mathbf{I}_\pm are the identities associated with $\mathfrak{u}(N-k)$ and $\mathfrak{u}(k)$, respectively:

$$\mathbf{I}_\pm = \frac{1}{2} (\mathbf{I} \pm \mathbf{Z}_k). \quad (6.11)$$

These deformed $\mathfrak{gl}_2 \oplus \mathfrak{gl}_2$ generators satisfy also the same relation as (3.11), and they are related to \mathbf{M}_{ab} and \mathbf{P}_a (6.5) analogously to (3.14) by

$$\mathbf{M}_{ab} = \frac{1}{2} \epsilon_{ab}^c (\mathbf{J}_c + \tilde{\mathbf{J}}_c) \quad \text{and} \quad \mathbf{P}_a = \frac{1}{2\sqrt{\sigma}} (\mathbf{J}_a - \tilde{\mathbf{J}}_a). \quad (6.12)$$

Therefore, if we define the matter fields using \mathbf{J}_a and $\tilde{\mathbf{J}}_a$, then they will have the standard interactions with the gravity.

We now introduce each elements of (6.9). The first one \mathfrak{c}_{CM} is the residual color symmetry:

$$\mathfrak{c}_{\text{CM}} = \mathfrak{iso} \otimes \left(\mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1) \right), \quad (6.13)$$

describing colored spin-two fields associated with the one form

$$\mathcal{C}_{\text{CM}} = \frac{1}{\ell_k} \left[(\varphi_+^a + \varphi_-^a) \mathbf{J}_a + (\tilde{\varphi}_+^a + \tilde{\varphi}_-^a) \tilde{\mathbf{J}}_a + (\psi^a \mathbf{J}_a + \tilde{\psi}^a \tilde{\mathbf{J}}_a) \mathbf{Y}_k \mathbf{Z}_k \right]. \quad (6.14)$$

The fields φ_+^a and $\tilde{\varphi}_+^a$ take values in $\mathfrak{su}(N-k)$, whereas φ_-^a and $\tilde{\varphi}_-^a$ in $\mathfrak{su}(k)$, both transforming in the adjoint representations. The fields ψ^a and $\tilde{\psi}^a$ are charged under $\mathfrak{u}(1)$. The matrix factor $\mathbf{Y}_k \mathbf{Z}_k$ is inserted to ensure $\text{Tr}(\mathfrak{b}_{\text{GR}} \mathfrak{c}_{\text{CM}}) = 0$, equivalently,

$$\text{Tr}(\mathbf{J} \mathbf{Y}_k \mathbf{Z}_k) = 0 = \text{Tr}(\tilde{\mathbf{J}} \mathbf{Y}_k \mathbf{Z}_k). \quad (6.15)$$

The second element \mathfrak{c}_{BS} is what corresponds to the broken part of the color symmetries:

$$\mathfrak{c}_{\text{BS}} = (\text{id} \oplus \mathfrak{iso}) \otimes \mathfrak{bs}. \quad (6.16)$$

Unlike the fields in \mathcal{C}_{CM} , this part does not describe massless spin-two fields. Rather, it describes so-called partially-massless spin-two fields [31, 32], as we shall demonstrate in the following. The corresponding one form is given by

$$\mathcal{C}_{\text{BS}} = \frac{1}{\ell_k} \left(\phi \mathbf{J} + \phi^a \mathbf{J}_a + \tilde{\phi} \tilde{\mathbf{J}} + \tilde{\phi}^a \tilde{\mathbf{J}}_a \right), \quad (6.17)$$

where the fields ϕ^a , ϕ , $\tilde{\phi}^a$ and $\tilde{\phi}$ take values in \mathfrak{bs} , carrying the bi-fundamental representations of $\mathfrak{su}(N-k)$ and $\mathfrak{su}(k)$, as well as the representation of $\mathfrak{u}(1)$. Because these fields anti-commute with \mathbf{Z}_k , they also intertwine the left-moving and the right-moving \mathfrak{gl}_2 's. For instance,

$$\phi^a \mathbf{J}_b = \tilde{\mathbf{J}}_b \phi^a. \quad (6.18)$$

As a consequence, they transform differently under Hermitian conjugate:

$$(\phi, \phi^a, \tilde{\phi}, \tilde{\phi}^a)^\dagger = \begin{cases} (-\tilde{\phi}, \tilde{\phi}^a, -\phi, \phi^a) & [\sigma = +1] \\ (-\phi, \phi^a, -\tilde{\phi}, \tilde{\phi}^a) & [\sigma = -1] \end{cases}, \quad (6.19)$$

compared to the massless ones (3.28).

6.2 Colored gravity around non-singlet vacua

With the precise form of the fields (6.6), (6.14), (6.17), we now rewrite the Chern-Simons action into a metric form. It is given by the sum of three terms as in (4.3). Firstly, we have the standard gravity action

$$S_{\text{Gravity}} = \frac{1}{16\pi G} \int d^3x \sqrt{|g|} \left(R + \frac{2\sigma}{\ell_k^2} \right), \quad (6.20)$$

with a k -dependent cosmological constant, set by (6.7). The Chern-Simons action S_{CS} for the gauge fields \mathbf{A}_+ for $\mathfrak{su}(N-k)$, \mathbf{A}_- for $\mathfrak{su}(k)$ and A for $\mathfrak{u}(1)$ are given analogously to (4.6). Finally, the action for the matter sector takes the following form:

$$S_{\text{Matter}} = \frac{1}{16\pi G} \int \frac{1}{N-2k} \left(L[\varphi_+, \tilde{\varphi}_+, \ell_k] - L[\varphi_-, \tilde{\varphi}_-, \ell_k] + L_{\text{BS}}[\phi, \tilde{\phi}, \ell_k] + L_{\text{cross}} \right) - \frac{k(N-k)}{N^2} L[\psi, \tilde{\psi}, \ell_k] - \frac{1}{\ell_k^2} \epsilon_{abc} e^a \wedge \Omega^b \wedge \Omega^c, \quad (6.21)$$

where L is the massless Lagrangian given in (4.8) whereas L_{BS} is given by

$$L_{\text{BS}}[\phi, \tilde{\phi}, \ell] = \frac{4\sqrt{\sigma}}{\ell} \text{Tr} \left[\left\{ \tilde{\phi} \wedge \left(D\phi - \frac{1}{\sqrt{\sigma}\ell} e^a \wedge \phi_a \right) - \tilde{\phi}_a \wedge \left(D\phi^a - \frac{1}{\sqrt{\sigma}\ell} e^a \wedge \phi \right) \right\} \mathbf{Z}_k \right]. \quad (6.22)$$

The covariant derivatives $D\phi^a$ and $D\phi$ are given by

$$\begin{aligned} D\phi^a &= D_\omega \phi^a + \left(\tilde{\mathbf{A}}_+ + \tilde{\mathbf{A}}_- + \tilde{\mathbf{A}} \mathbf{Y}_k \right) \wedge \phi^a - \phi^a \wedge (\mathbf{A}_+ + \mathbf{A}_- + A \mathbf{Y}_k), \\ D\phi &= d\phi + \left(\tilde{\mathbf{A}}_+ + \tilde{\mathbf{A}}_- + \tilde{\mathbf{A}} \mathbf{Y}_k \right) \wedge \phi - \phi \wedge (\mathbf{A}_+ + \mathbf{A}_- + A \mathbf{Y}_k), \end{aligned} \quad (6.23)$$

and similarly for the tilde counter parts. The other terms in (6.21) give additional interactions: the last term gives quartic interaction through $\Omega_a{}^b$:

$$\Omega_a{}^b = \frac{1}{N-2k} \left[W_b^a(\varphi_+, \tilde{\varphi}_+) + W_b^a(\varphi_-, \tilde{\varphi}_-) + W_{BSb}^a(\phi, \tilde{\phi}) \right] - \frac{1}{k(N-k)} W_b^a(\psi, \tilde{\psi}), \quad (6.24)$$

where W_b^a is given by (4.12) and W_{BSb}^a by

$$W_{BSb}^a(\phi, \tilde{\phi}) = 8\sqrt{\sigma} \text{Tr} \left[\left(\tilde{\phi}_{[a}{}^{[b} \phi_{c]}{}^{c]} - \frac{1}{4} \delta_a^b \tilde{\phi}_{[c}{}^c \phi_{d]}{}^d \right) \mathbf{Z}_k \right]. \quad (6.25)$$

The term L_{cross} , given by

$$L_{\text{cross}} = \frac{4\sqrt{\sigma}}{\ell_k^2} \text{Tr} \left[(\varphi_+^a - \varphi_-^a + \psi^a \mathbf{Y}_k) \wedge (\tilde{\phi} \wedge \phi_a + \tilde{\phi}_a \wedge \phi) + \right. \\ \left. + \epsilon_{abc} (\varphi_+^a \wedge \varphi_+^b + \varphi_-^a \wedge \varphi_-^b) \wedge \psi^c \mathbf{Y}_k \mathbf{Z}_k \right] - ([] \leftrightarrow [\tilde{\ }]), \quad (6.26)$$

is the cross terms originating from the Chern-Simons cubic interactions.

In principle, we can further simplify the action as we did in the singlet vacuum case. However, already at this level, we can extract a lot of physics.

- We have a scalar potential as a function of four fields φ_{\pm} , ψ , ϕ (and their tilde counter parts) and the point where all fields vanish correspond to the extremum point whose potential value gives the cosmological constant $-\sigma/\ell_k^2$. This potential should be a shift of the potential $V(\varphi, \tilde{\varphi})$ (4.15) defined around the singlet vacuum, hence it will admit all other vacua as extrema.
- The interaction strength for each field can be easily read off from the action. The gravity and gauge interaction have the same strength controlled by G and κ as in the singlet vacuum case. The interaction of colored spin two fields φ_{\pm} is weakened — the coefficient changed from N to $N-2k$. The same for the broken-symmetry field ϕ . Finally, ψ has interaction strength controlled by $N^2/[k(N-k)]$. Therefore, when the color symmetry is maximally broken, that is $N-2k \sim 1$, the interaction between all these fields becomes as weak as the gravitational one.
- Let us conclude this section with the summary of the field content around the k -vacuum. At first, we have the graviton and $\mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1)$ Chern-Simons gauge fields. Next, about the colored matter fields, there are $(N-k)^2 - 1$ fields for $(\varphi_+, \tilde{\varphi}_+)$, $k^2 - 1$ for $(\varphi_-, \tilde{\varphi}_-)$ and 1 for $(\psi, \tilde{\psi})$. They are all massless spin-two fields, but $(\varphi_-, \tilde{\varphi}_-)$ and $(\psi, \tilde{\psi})$ — hence k^2 fields — are in fact ghost. For the broken symmetry part, we have $2k(N-k)$ fields for $(\phi, \tilde{\phi})$. The latter describes so-called partially-massless fields and its proper analysis is the subject of the next section.

6.3 Partially massless spectrum associated with broken color symmetry

Around a non-singlet vacuum, the fields φ_{\pm} and ψ both describe massless spin-two fields having the same quadratic Lagrangian given by (4.8). On the other hand, the fields ϕ

corresponding to the broken part of the color symmetries have different quadratic Lagrangian (6.22), hence describe different spectrum. We have already mentioned that they correspond to partially-massless fields [31, 32]. In this section, we analyze the quadratic Lagrangian (6.22) to prove this statement. Here, we concentrate on AdS_3 . To get the dS_3 result, it is sufficient to replace ℓ by $i\ell$.

Though the Lagrangian (6.22) has a rather non-standard form involving cross term between ϕ and $\tilde{\phi}$ together with an insertion of \mathbf{Z}_k , it can always be diagonalized with the help of the Hermiticity property (6.19). Therefore, for the spectrum analysis, it will suffice to consider $S_{\text{BS}}[\phi, \phi^a]$ taking the following expression:

$$S_{\text{BS}}[\phi, \phi^a] = \int \phi \wedge \left(d\phi - \frac{1}{\ell} e_a \wedge \phi^a \right) - \phi_a \wedge \left(D\phi^a - \frac{1}{\ell} e^a \wedge \phi \right), \quad (6.27)$$

with the AdS dreibein and spin connection (e^a, ω^{ab}) . We first note that this action admits the gauge symmetries with parameters $(\varepsilon, \varepsilon^a)$,

$$\delta \phi = d\varepsilon - \frac{1}{\ell} e^a \varepsilon_a, \quad \delta \phi^a = D\varepsilon^a - \frac{1}{\ell} e^a \varepsilon, \quad (6.28)$$

which come from the Chern-Simons gauge symmetries.

For a closer look of this action involving three fields $h_{\mu\nu} = e^a_{(\mu} \phi_{\nu)a}$, $f_{\mu\nu} = e^a_{[\mu} \phi_{\nu]a}$ and $\phi_\mu = e^a_\mu \phi_a$, we consider two different but equivalent paths:

- We first derive the equation of motion for one-form fields ϕ^a and ϕ . They are given by

$$D\phi^a - \frac{1}{\ell} e^a \wedge \phi = 0, \quad d\phi - \frac{1}{\ell} e^a \wedge \phi_a = 0. \quad (6.29)$$

The second equation implies that the antisymmetric field $f_{\mu\nu}$ is the field strength of ϕ_μ : $f_{\mu\nu} = \ell \partial_{[\mu} \phi_{\nu]}$. Then, by gauge fixing ϕ_μ to zero with the gauge parameter ε^a , the field $f_{\mu\nu}$ decouples from the first equation. We thus end up with only one field $h_{\mu\nu}$ satisfying the equation of motion,

$$\nabla_{[\mu} h_{\nu]\rho} = 0, \quad (6.30)$$

and the gauge symmetry,

$$\delta h_{\mu\nu} = \left(\nabla_\mu \nabla_\nu - \frac{1}{\ell^2} g_{\mu\nu} \right) \varepsilon. \quad (6.31)$$

This coincides with the gauge symmetry of partially-massless spin-two field [31, 32].

- Instead of first deriving the equation and then gauge fixing to $\phi_\mu = 0$, one can reverse the procedure. We first gauge fix and eliminate ϕ_μ field in the action and obtain

$$S_{\text{PM}}[\phi_{\mu\nu}] = \int d^3x \sqrt{|g|} \epsilon^{\mu\nu\rho} \phi^\lambda_{\mu} \nabla_\nu \phi_{\rho\lambda}, \quad (6.32)$$

modulo a boundary term. We note that the field $\phi_{\mu\nu}$ contains both of the symmetric part $h_{\mu\nu}$ and the antisymmetric part $f_{\mu\nu}$. Only $h_{\mu\nu}$ admits the gauge symmetries (6.31). The equation of motion is now given by

$$C_{\mu\nu,\rho} := \nabla_{[\mu} h_{\nu]\rho} + \nabla_{[\mu} f_{\nu]\rho} = 0. \quad (6.33)$$

The totally anti-symmetric part $C_{[\mu\nu,\rho]} = \partial_{[\mu} f_{\nu\rho]} = 0$ can be readily solved as $f_{\mu\nu} = \partial_{[\mu} a_{\nu]}$. With the field redefinition $h_{\mu\nu} \rightarrow h_{\mu\nu} - \nabla_{(\mu} a_{\nu)}$, the trace of the above equation, $C^\rho_{\mu,\rho} = 0$, gives

$$a_\mu = \frac{\ell^2}{2} (\nabla^\rho h_{\mu\rho} - \nabla_\mu h^\rho{}_\rho), \quad (6.34)$$

Taking now a divergence of $C_{\mu\nu,\rho}$, we arrive at the second-order equation,

$$\nabla^\rho C_{\rho(\mu,\nu)} = G_{\mu\nu}^{\text{lin}} + \frac{1}{\ell^2} (h_{\mu\nu} - g_{\mu\nu} h^\rho{}_\rho) = 0, \quad (6.35)$$

with the linearized Einstein tensor $G_{\mu\nu}^{\text{lin}}$. One can also check that the mass in the above equation corresponds to that of a partially-massless field. Furthermore, using Bianchi identity, we deduce that the left-hand side of (6.34) vanishes, so does $f_{\mu\nu}$. Therefore, we end up with the same equation (6.30).⁹

7 Discussions

In this paper, we proposed a Chan-Paton color-decorated gravity in three dimensions and studied its properties. We have shown that the theory describes a gravitational system of colored massless spin-two matter fields coupled to $\mathfrak{su}(N)$ gauge fields. These matter fields have a non-trivial potential whose extrema have $\lfloor \frac{N+1}{2} \rfloor$ different values of cosmological constant. All the extremum points but the origin spontaneously break the $\mathfrak{su}(N)$ color symmetry down to $\mathfrak{su}(N-k) \oplus \mathfrak{su}(k) \oplus \mathfrak{u}(1)$. We found that the spin-two Goldstone modes corresponding to the broken part of the symmetries are combined with the gauge fields and become partially-massless spin-two fields. In the vacua with large $k \sim N/2$, the interactions of the matter fields are as weak as the gravitational one. In the small k vacua, their interaction becomes strong by the factor of \sqrt{N} .

Considering the dS_3 branch, the potential takes a spiral stairwell shape (figure 3) with $\lfloor \frac{N+1}{2} \rfloor$ many steps, having split cosmological constants that range from $\Lambda = 1/\ell^2$ at the lowest step all the way up to $\sim N^2 \Lambda$ at the highest step. The spacing gets dense in lower steps, while sparse in higher steps. If such features continue to hold in higher dimensions, the colored gravity with large N might be very relevant for the early universe cosmology in that the universe begins in an inflationary epoch with a large cosmological constant at a very high stairstep. The colored matter are weakly coupled there, and hence they are not confined. As the state of the universe decays towards lower stairsteps, the effective cosmological constant decreases sequentially and eventually exits the inflation. The colored matter fields start to interact stronger and eventually form heavy color-neutral composites. It is in this synopsis that the spin-two colored matter fields might play a novel role in the current paradigm of the inflationary cosmology.

⁹Strictly speaking, the equation (6.35) alone is weaker than the first-order one (6.30). The former describes one propagating degrees of freedom, while the latter does not have any bulk mode and corresponds to the spectrum described by (6.27). Note that the latter partially-massless spectrum is what the three-dimensional conformal gravity contains analogously to the four-dimensional case [45–47]. To recapitulate, in three dimensions (not in higher dimensions), there are two kinds of partially-massless fields for the maximal depth, which includes the spin-two partially-massless spectrum. We shall discuss more about this subtlety in the companion paper [27].

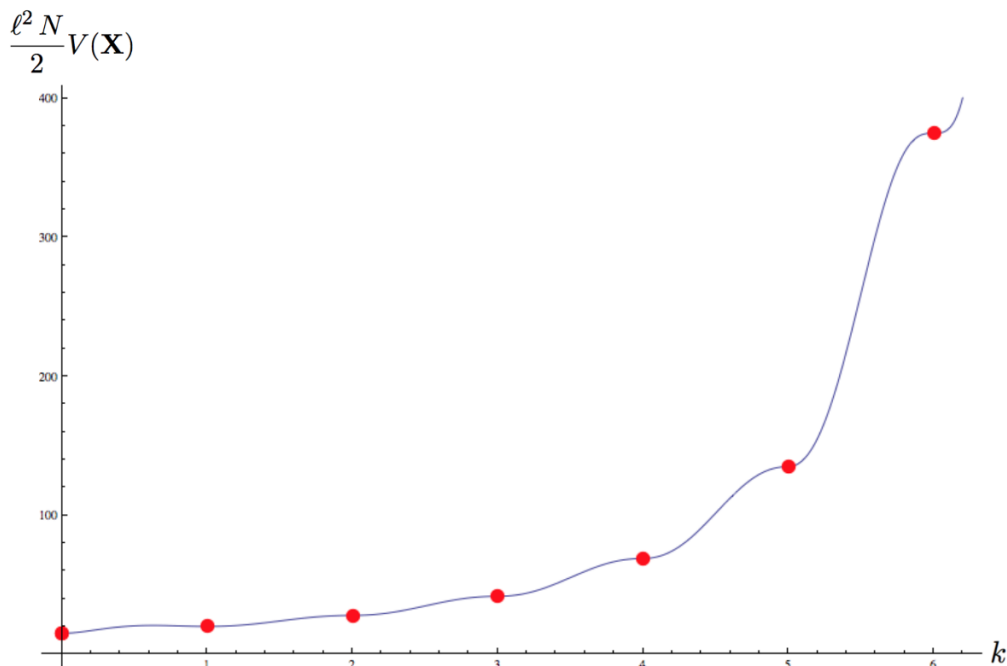


Figure 3. Potential of the colored gravity in dS ($N = 15$): k is the parameter of a curve in $\mathfrak{su}(15)$ that passes through the extremum points.

We also speculate on a novel approach to the three-dimensional quantum colored gravity. At large N , the contribution of the $\mathcal{O}(N/2)$ multiple vacua in the path integral might be captured by the $\mathfrak{su}(N)$ random matrix model given by

$$\mathcal{Z}_{\text{MM}} = \int d\mathbf{X} \exp[i c V(\mathbf{X})]. \quad (7.1)$$

It would be also interesting to explore ab initio definition of the three-dimensional quantum gravity starting from tensor-field valued matrix models.

This work brings in many open problems worth of further investigation. First of all, extensions to (higher-spin) supergravity as well as the analysis of the asymptotic symmetries [48–50] are imminent. Further extensions to color-decoration of the known higher-spin gravity in three-dimensional Lifshitz spacetime [51] and flat spacetime [52, 53] are also straightforward. Extension to higher-dimensional spacetime is also highly interesting. A version of such situation was already studied in the context of AdS/CFT correspondence [54]. Vasiliev equations for color-decorated higher-spin theories needs to be better understood, along with higher-dimensional counterpart of the staircase potential we found in three dimensions. As the color dynamics is described by Chern-Simons gauge theory, one might anticipate to formulate colored gravity in any dimensions in terms of a version of Chern-Simons formulation, perhaps, along the lines of [55] and [56]. Quantum aspects of color-decorated gravity is an avenue to be explored. In particular, consequences and implications of strong color interactions among colored spin-two fields. Turning to the inflationary cosmology, it would be interesting to understand how the color-decoration mod-

ifies the infrared dynamics of interacting massless spin-two fields at super-horizon scales. This brings one to investigate stochastic dynamics of these fields, as would be described by color-decorated version of the Langevin dynamics [57, 58].

Acknowledgments

We are grateful to Marc Henneaux, Jaewon Kim, Jihun Kim, Sasha Polyakov, Augusto Sagnotti and Misha Vasiliev for many useful discussions. This work was supported in part by the National Research Foundation of Korea through the grant NRF-2014R1A6A3A04056670 (SG, EJ), and the grants 2005-0093843, 2010-220-C00003 and 2012K2A1A9055280 (SG, KM, SJR). The work of EJ is also supported by the Russian Science Foundation grant 14-42-00047 associated with Lebedev Institute. The work of KM was supported by the BK21 Plus Program funded by the Ministry of Education(MOE, Korea) and National Research Foundation of Korea (NRF).

A Spin-two fields in three dimensions

In this section, we briefly review massive and (partially) massless spin-two field in three dimensions. Let us begin with the standard Fierz-Pauli massive spin-two action in (A)dS_{d+1} background,

$$\begin{aligned} \mathcal{L}_{\text{FP}}(\chi) = & -\frac{1}{2} \nabla^\rho \chi^{\mu\nu} \nabla_\rho \chi_{\mu\nu} + \nabla^\nu \chi_{\mu\nu} \nabla_\rho \chi^{\mu\rho} - \nabla_\mu \chi^\rho{}_\rho \nabla_\nu \chi^{\mu\nu} + \frac{1}{2} \nabla_\mu \chi^\rho{}_\rho \nabla^\mu \chi^\sigma{}_\sigma \\ & + \frac{\sigma}{\ell^2} \left(\chi^{\mu\nu} \chi_{\mu\nu} + \frac{d-2}{2} \chi^\mu{}_\mu \chi^\nu{}_\nu \right) - \frac{1}{2} m^2 (\chi^{\mu\nu} \chi_{\mu\nu} - \chi^\mu{}_\mu \chi^\nu{}_\nu), \end{aligned} \quad (\text{A.1})$$

whose equations of motion reduce to the Fierz system:

$$\left(\nabla^2 + \frac{2\sigma}{\ell^2} - m^2 \right) \chi_{\mu\nu} = 0, \quad \chi^\mu{}_\mu = 0 = \nabla^\mu \chi_{\mu\nu}. \quad (\text{A.2})$$

In terms of the lowest energy Δ , the mass parameter m is given as

$$\sigma \ell^2 m^2 = \Delta (\Delta - d) = \mu^2 - \frac{d^2}{4}, \quad \mu = \Delta - \frac{d}{2}. \quad (\text{A.3})$$

In AdS, a very massive field corresponds to a large real μ , whereas in dS it corresponds to a large pure imaginary μ . When the mass term of the action takes a special value, the action acquires gauge symmetries: there are two such points, $m^2 = 0$ for *massless-ness* and $m^2 = -\frac{d-1}{\ell^2} \sigma$ for *partially-massless-ness*. Due to the gauge symmetries, these spectra have smaller number of DoF than the massive one. In particular, partially-massless spin-two field has the same amount of DoF as massless spin-two and massless spin-one fields. The basic properties of massless and partially-massless spectra are summarized in table 1.

In three dimensions, any spin-two spectrum can be described in terms of a first-derivative Lagrangian. Again beginning with the massive Lagrangian (A.1), we can reformulate the Lagrangian into

$$\mathcal{L}_{\text{FD}}(\chi, \tau) = \frac{1}{2} \epsilon^{\mu\nu\rho} \left(\tau_\mu{}^\lambda \nabla_\nu \chi_{\rho\lambda} + \chi_\mu{}^\lambda \nabla_\nu \tau_{\rho\lambda} \right) + \mu \left(\tau_{[\mu}{}^\mu \tau_{\nu]}{}^\nu + \frac{\sigma}{\ell^2} \chi_{[\mu}{}^\mu \chi_{\nu]}{}^\nu \right), \quad (\text{A.4})$$

Spectrum	m^2	Δ_+	μ	Gauge symmetry
Massless	0	d	$\frac{d}{2}$	$\delta \chi_{\mu\nu} = \nabla_{(\mu} \xi_{\nu)}$
partially-massless	$-\frac{d-1}{\ell^2} \sigma$	$d-1$	$\frac{d-2}{2}$	$\delta \chi_{\mu\nu} = (\nabla_\mu \nabla_\nu - \frac{\sigma}{\ell^2} g_{\mu\nu}) \xi$

Table 1. Massless and partially-massless spin-two fields.

by introducing an auxiliary field $\tau_{\mu\nu}$. Here, the tensors $\chi_{\mu\nu}$ and $\tau_{\mu\nu}$ do not have any symmetry properties. By integrating out $\tau_{\mu\nu}$ — that is by plugging in the solution $\tau_{\mu\nu}(\chi)$ of its own equation — one can show that the antisymmetric part of $\tau_{\mu\nu}$ drops and the Lagrangian (A.4) reproduces the Fierz-Pauli Lagrangian (A.1) up to a factor :

$$\mathcal{L}_{\text{FD}}(\chi, \tau(\chi)) = \frac{1}{\mu} \mathcal{L}_{\text{FP}}(\chi). \quad (\text{A.5})$$

It is more convenient to recast the Lagrangian (A.4) in terms of $\varphi_{\mu\nu}$ and $\tilde{\varphi}_{\mu\nu}$:

$$\chi_{\mu\nu} = \sqrt{\sigma} \ell (\varphi_{\mu\nu} - \tilde{\varphi}_{\mu\nu}), \quad \tau_{\mu\nu} = \varphi_{\mu\nu} + \tilde{\varphi}_{\mu\nu}. \quad (\text{A.6})$$

so that the massive spin-two Lagrangian splits into the parity breaking spin +2 and spin −2 parts:

$$\mathcal{L}_{\text{FD}}(\chi, \tau) = \mathcal{L}_{+2}(\varphi) + \mathcal{L}_{-2}(\tilde{\varphi}). \quad (\text{A.7})$$

Here, the self-dual massive spin ± 2 Lagrangian [59] is given by

$$\mathcal{L}_{\pm 2}(\varphi) = \pm \sqrt{\sigma} \ell \epsilon^{\mu\nu\rho} \varphi_\mu{}^\lambda \nabla_\nu \varphi_{\rho\lambda} + 2 \mu \varphi_{[\mu}{}^\mu \varphi_{\nu]}{}^\nu. \quad (\text{A.8})$$

Let us remark an unusual feature of this parity breaking massive spin-two Lagrangian in three dimensions: the sign of the *mass-like* term actually determines whether the Lagrangian is ghost or not, the positive sign for the unitary case and the negative sign for the ghost, whereas the sign of the *kinetic-like* term determines the sign of the spin. These can be seen, for example, by dualizing the above first-derivative Lagrangian to the second-order Lagrangian and relating μ to the sign in front of this Lagrangian. For this reason, one can render a unitary spin-two field to a ghost one by only modifying its mass-like term in the first order description of three dimensional theories. Throughout this paper, we encounter three different cases: firstly, the $\mu = 1$ case corresponds to unitary massless spin-two field, whereas the $\mu = -1$ case gives ghost massless spin-two field. The case of $\mu = 0$ describes partially-massless spin-two field, which does not admit any two-derivative description as is clear from (A.4) and (A.5).

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References

- [1] T. Clifton, P.G. Ferreira, A. Padilla and C. Skordis, *Modified Gravity and Cosmology*, *Phys. Rept.* **513** (2012) 1 [[arXiv:1106.2476](https://arxiv.org/abs/1106.2476)] [[INSPIRE](#)].
- [2] A. De Felice and S. Tsujikawa, *$f(R)$ theories*, *Living Rev. Rel.* **13** (2010) 3 [[arXiv:1002.4928](https://arxiv.org/abs/1002.4928)] [[INSPIRE](#)].
- [3] S. Capozziello and M. De Laurentis, *Extended Theories of Gravity*, *Phys. Rept.* **509** (2011) 167 [[arXiv:1108.6266](https://arxiv.org/abs/1108.6266)] [[INSPIRE](#)].
- [4] D.G. Boulware and S. Deser, *Can gravitation have a finite range?*, *Phys. Rev. D* **6** (1972) 3368 [[INSPIRE](#)].
- [5] P. Creminelli, A. Nicolis, M. Papucci and E. Trincherini, *Ghosts in massive gravity*, *JHEP* **09** (2005) 003 [[hep-th/0505147](https://arxiv.org/abs/hep-th/0505147)] [[INSPIRE](#)].
- [6] E. Joung, W. Li and M. Taronna, *No-Go Theorems for Unitary and Interacting Partially Massless Spin-Two Fields*, *Phys. Rev. Lett.* **113** (2014) 091101 [[arXiv:1406.2335](https://arxiv.org/abs/1406.2335)] [[INSPIRE](#)].
- [7] C. de Rham and G. Gabadadze, *Generalization of the Fierz-Pauli Action*, *Phys. Rev. D* **82** (2010) 044020 [[arXiv:1007.0443](https://arxiv.org/abs/1007.0443)] [[INSPIRE](#)].
- [8] C. de Rham, G. Gabadadze and A.J. Tolley, *Resummation of Massive Gravity*, *Phys. Rev. Lett.* **106** (2011) 231101 [[arXiv:1011.1232](https://arxiv.org/abs/1011.1232)] [[INSPIRE](#)].
- [9] S.F. Hassan and R.A. Rosen, *Resolving the Ghost Problem in non-Linear Massive Gravity*, *Phys. Rev. Lett.* **108** (2012) 041101 [[arXiv:1106.3344](https://arxiv.org/abs/1106.3344)] [[INSPIRE](#)].
- [10] S.F. Hassan and R.A. Rosen, *Confirmation of the Secondary Constraint and Absence of Ghost in Massive Gravity and Bimetric Gravity*, *JHEP* **04** (2012) 123 [[arXiv:1111.2070](https://arxiv.org/abs/1111.2070)] [[INSPIRE](#)].
- [11] E.A. Bergshoeff, O. Hohm and P.K. Townsend, *Massive Gravity in Three Dimensions*, *Phys. Rev. Lett.* **102** (2009) 201301 [[arXiv:0901.1766](https://arxiv.org/abs/0901.1766)] [[INSPIRE](#)].
- [12] E. Bergshoeff, W. Merbis, A.J. Routh and P.K. Townsend, *The Third Way to 3D Gravity*, *Int. J. Mod. Phys. D* **24** (2015) 1544015 [[arXiv:1506.05949](https://arxiv.org/abs/1506.05949)] [[INSPIRE](#)].
- [13] X.O. Camanho, J.D. Edelstein, J. Maldacena and A. Zhiboedov, *Causality Constraints on Corrections to the Graviton Three-Point Coupling*, *JHEP* **02** (2016) 020 [[arXiv:1407.5597](https://arxiv.org/abs/1407.5597)] [[INSPIRE](#)].
- [14] H. Lü, A. Perkins, C.N. Pope and K.S. Stelle, *Black Holes in Higher-Derivative Gravity*, *Phys. Rev. Lett.* **114** (2015) 171601 [[arXiv:1502.01028](https://arxiv.org/abs/1502.01028)] [[INSPIRE](#)].
- [15] R.M. Wald, *Spin-2 Fields and General Covariance*, *Phys. Rev. D* **33** (1986) 3613 [[INSPIRE](#)].
- [16] C. Cutler and R.M. Wald, *A New Type of Gauge Invariance for a Collection of Massless Spin-2 Fields. 1. Existence and Uniqueness*, *Class. Quant. Grav.* **4** (1987) 1267 [[INSPIRE](#)].
- [17] R.M. Wald, *A New Type of Gauge Invariance for a Collection of Massless Spin-2 Fields. 2. Geometrical Interpretation*, *Class. Quant. Grav.* **4** (1987) 1279 [[INSPIRE](#)].

- [18] N. Boulanger, T. Damour, L. Gualtieri and M. Henneaux, *Inconsistency of interacting, multigraviton theories*, *Nucl. Phys. B* **597** (2001) 127 [[hep-th/0007220](#)] [[INSPIRE](#)].
- [19] M.A. Vasiliev, *Consistent equation for interacting gauge fields of all spins in (3+1)-dimensions*, *Phys. Lett. B* **243** (1990) 378 [[INSPIRE](#)].
- [20] S.F. Prokushkin and M.A. Vasiliev, *Higher spin gauge interactions for massive matter fields in 3-D AdS space-time*, *Nucl. Phys. B* **545** (1999) 385 [[hep-th/9806236](#)] [[INSPIRE](#)].
- [21] M.A. Vasiliev, *Nonlinear equations for symmetric massless higher spin fields in (A)dS₄*, *Phys. Lett. B* **567** (2003) 139 [[hep-th/0304049](#)] [[INSPIRE](#)].
- [22] M.A. Vasiliev, *Extended Higher Spin Superalgebras and Their Realizations in Terms of Quantum Operators*, *Fortsch. Phys.* **36** (1988) 33 [[INSPIRE](#)].
- [23] S.E. Konstein and M.A. Vasiliev, *Massless Representations and Admissibility Condition for Higher Spin Superalgebras*, *Nucl. Phys. B* **312** (1989) 402 [[INSPIRE](#)].
- [24] M.A. Vasiliev, *Consistent Equations for Interacting Massless Fields of All Spins in the First Order in Curvatures*, *Annals Phys.* **190** (1989) 59 [[INSPIRE](#)].
- [25] S.E. Konstein and M.A. Vasiliev, *Extended Higher Spin Superalgebras and Their Massless Representations*, *Nucl. Phys. B* **331** (1990) 475 [[INSPIRE](#)].
- [26] M.A. Vasiliev, *Higher spin superalgebras in any dimension and their representations*, *JHEP* **12** (2004) 046 [[hep-th/0404124](#)] [[INSPIRE](#)].
- [27] S. Gwak, E. Joung, K. Mkrtchyan and S.-J. Rey, *Rainbow Vacua of Colored Higher Spin Gravity in Three Dimensions*, [arXiv:1511.05975](#) [[INSPIRE](#)].
- [28] A. Achucarro and P.K. Townsend, *A Chern-Simons Action for Three-Dimensional anti-de Sitter Supergravity Theories*, *Phys. Lett. B* **180** (1986) 89 [[INSPIRE](#)].
- [29] E. Witten, *(2 + 1)-Dimensional Gravity as an Exactly Soluble System*, *Nucl. Phys. B* **311** (1988) 46 [[INSPIRE](#)].
- [30] M. Bañados, C. Teitelboim and J. Zanelli, *The black hole in three-dimensional space-time*, *Phys. Rev. Lett.* **69** (1992) 1849 [[hep-th/9204099](#)] [[INSPIRE](#)].
- [31] S. Deser and A. Waldron, *Gauge invariances and phases of massive higher spins in (A)dS*, *Phys. Rev. Lett.* **87** (2001) 031601 [[hep-th/0102166](#)] [[INSPIRE](#)].
- [32] S. Deser and A. Waldron, *Null propagation of partially massless higher spins in (A)dS and cosmological constant speculations*, *Phys. Lett. B* **513** (2001) 137 [[hep-th/0105181](#)] [[INSPIRE](#)].
- [33] S.N. Gupta, *Quantization of Einstein's gravitational field: general treatment*, *Proc. Phys. Soc. A* **65** (1952) 608 [[INSPIRE](#)].
- [34] R.H. Kraichnan, *Special-Relativistic Derivation of Generally Covariant Gravitation Theory*, *Phys. Rev.* **98** (1955) 1118 [[INSPIRE](#)].
- [35] V.I. Ogievetsky and I.V. Polubarinov, *Interacting Field of Spin 2 and the Einstein Equations*, *Ann. Phys.* **35** (1965) 167.
- [36] S. Deser, *Selfinteraction and gauge invariance*, *Gen. Rel. Grav.* **1** (1970) 9 [[gr-qc/0411023](#)] [[INSPIRE](#)].
- [37] R.P. Feynman, F.B. Morinigo, W.G. Wagner and B. Hatfield, *Feynman Lectures on Gravitation*, California Institute of Technology (1963), Addison-Wesley, Reading, U.S.A. (1995), pg. 232.

- [38] N. Boulanger and L. Gualtieri, *An Exotic theory of massless spin two fields in three-dimensions*, *Class. Quant. Grav.* **18** (2001) 1485 [[hep-th/0012003](#)] [[INSPIRE](#)].
- [39] S.C. Anco, *Parity violating spin-two gauge theories*, *Phys. Rev. D* **67** (2003) 124007 [[gr-qc/0305026](#)] [[INSPIRE](#)].
- [40] C. Aragone and S. Deser, *Consistency Problems of Hypergravity*, *Phys. Lett. B* **86** (1979) 161 [[INSPIRE](#)].
- [41] C. Aragone and S. Deser, *Consistency Problems of Spin-2 Gravity Coupling*, *Nuovo Cim. B* **57** (1980) 33 [[INSPIRE](#)].
- [42] E.S. Fradkin and M.A. Vasiliev, *On the Gravitational Interaction of Massless Higher Spin Fields*, *Phys. Lett. B* **189** (1987) 89 [[INSPIRE](#)].
- [43] E.S. Fradkin and M.A. Vasiliev, *Cubic Interaction in Extended Theories of Massless Higher Spin Fields*, *Nucl. Phys. B* **291** (1987) 141 [[INSPIRE](#)].
- [44] E. Joung and M. Taronna, *Cubic-interaction-induced deformations of higher-spin symmetries*, *JHEP* **03** (2014) 103 [[arXiv:1311.0242](#)] [[INSPIRE](#)].
- [45] S. Deser, E. Joung and A. Waldron, *Gravitational- and Self-Coupling of Partially Massless Spin 2*, *Phys. Rev. D* **86** (2012) 104004 [[arXiv:1301.4181](#)] [[INSPIRE](#)].
- [46] S. Deser, E. Joung and A. Waldron, *Partial Masslessness and Conformal Gravity*, *J. Phys. A* **46** (2013) 214019 [[arXiv:1208.1307](#)] [[INSPIRE](#)].
- [47] E. Joung and K. Mkrtchyan, *Partially-massless higher-spin algebras and their finite-dimensional truncations*, *JHEP* **01** (2016) 003 [[arXiv:1508.07332](#)] [[INSPIRE](#)].
- [48] M. Henneaux and S.-J. Rey, *Nonlinear W_∞ as Asymptotic Symmetry of Three-Dimensional Higher Spin Anti-de Sitter Gravity*, *JHEP* **12** (2010) 007 [[arXiv:1008.4579](#)] [[INSPIRE](#)].
- [49] A. Campoleoni, S. Fredenhagen, S. Pfenninger and S. Theisen, *Asymptotic symmetries of three-dimensional gravity coupled to higher-spin fields*, *JHEP* **11** (2010) 007 [[arXiv:1008.4744](#)] [[INSPIRE](#)].
- [50] M. Henneaux, G. Lucena Gómez, J. Park and S.-J. Rey, *Super- W_∞ Asymptotic Symmetry of Higher-Spin AdS_3 Supergravity*, *JHEP* **06** (2012) 037 [[arXiv:1203.5152](#)] [[INSPIRE](#)].
- [51] M. Gary, D. Grumiller, S. Prohazka and S.-J. Rey, *Lifshitz Holography with Isotropic Scale Invariance*, *JHEP* **08** (2014) 001 [[arXiv:1406.1468](#)] [[INSPIRE](#)].
- [52] H. Afshar, A. Bagchi, R. Fareghbal, D. Grumiller and J. Rosseel, *Spin-3 Gravity in Three-Dimensional Flat Space*, *Phys. Rev. Lett.* **111** (2013) 121603 [[arXiv:1307.4768](#)] [[INSPIRE](#)].
- [53] H.A. Gonzalez, J. Matulich, M. Pino and R. Troncoso, *Asymptotically flat spacetimes in three-dimensional higher spin gravity*, *JHEP* **09** (2013) 016 [[arXiv:1307.5651](#)] [[INSPIRE](#)].
- [54] O. Aharony, M. Berkooz and S.-J. Rey, *Rigid holography and six-dimensional $\mathcal{N} = (2, 0)$ theories on $AdS_5 \times S^1$* , *JHEP* **03** (2015) 121 [[arXiv:1501.02904](#)] [[INSPIRE](#)].
- [55] I. Bars and S.-J. Rey, *Noncommutative $Sp(2, \mathbb{R})$ gauge theories: A Field theory approach to two time physics*, *Phys. Rev. D* **64** (2001) 046005 [[hep-th/0104135](#)] [[INSPIRE](#)].
- [56] R. Bonezzi, O. Corradini, E. Latini and A. Waldron, *Quantum Gravity and Causal Structures: Second Quantization of Conformal Dirac Algebras*, *Phys. Rev. D* **91** (2015) 121501 [[arXiv:1505.01013](#)] [[INSPIRE](#)].

- [57] A.A. Starobinsky, *Stochastic de Sitter (inflationary) stage in the early universe*, *Lect. Notes Phys.* **246** (1986) 107 [[INSPIRE](#)].
- [58] S.-J. Rey, *Dynamics of Inflationary Phase Transition*, *Nucl. Phys. B* **284** (1987) 706 [[INSPIRE](#)].
- [59] C. Aragone and A. Khoudeir, *Selfdual massive gravity*, *Phys. Lett. B* **173** (1986) 141 [[INSPIRE](#)].